Orion Nebula, Distance about
450 pc, Total Mass about
5000 M(sun), Diameter about 3 pc

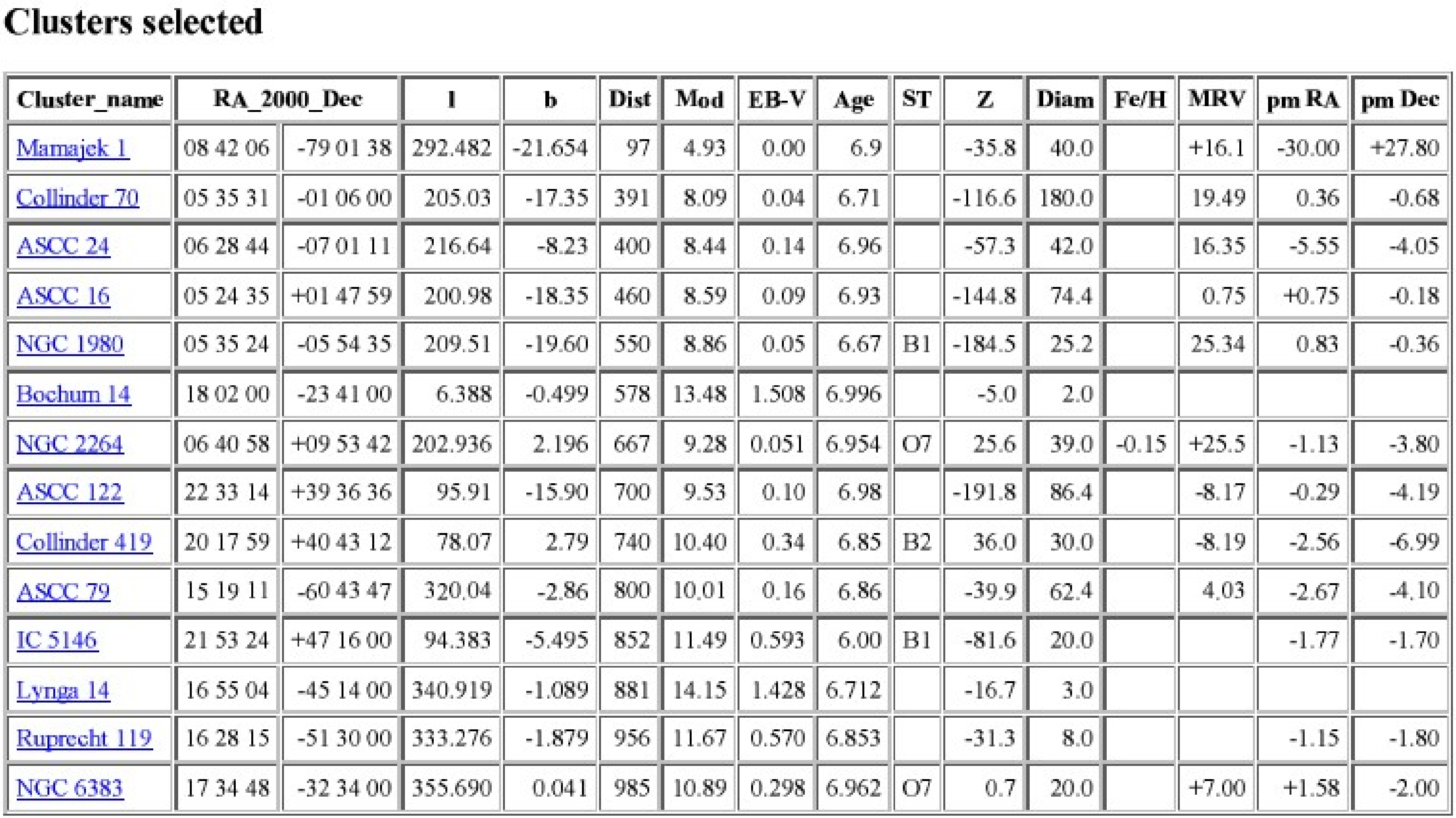
M11, NGC 6705: Total Mass
About 10000 M(sun), 200 Myr

Cluster formation

- Observations versus Models
- Important parameters
  1. Time scale
  2. Total mass
  3. Initial Mass Function
  4. Velocity distribution
  5. Binary fraction
  6. Diameter
  7. Density distribution

Heuristic Approach

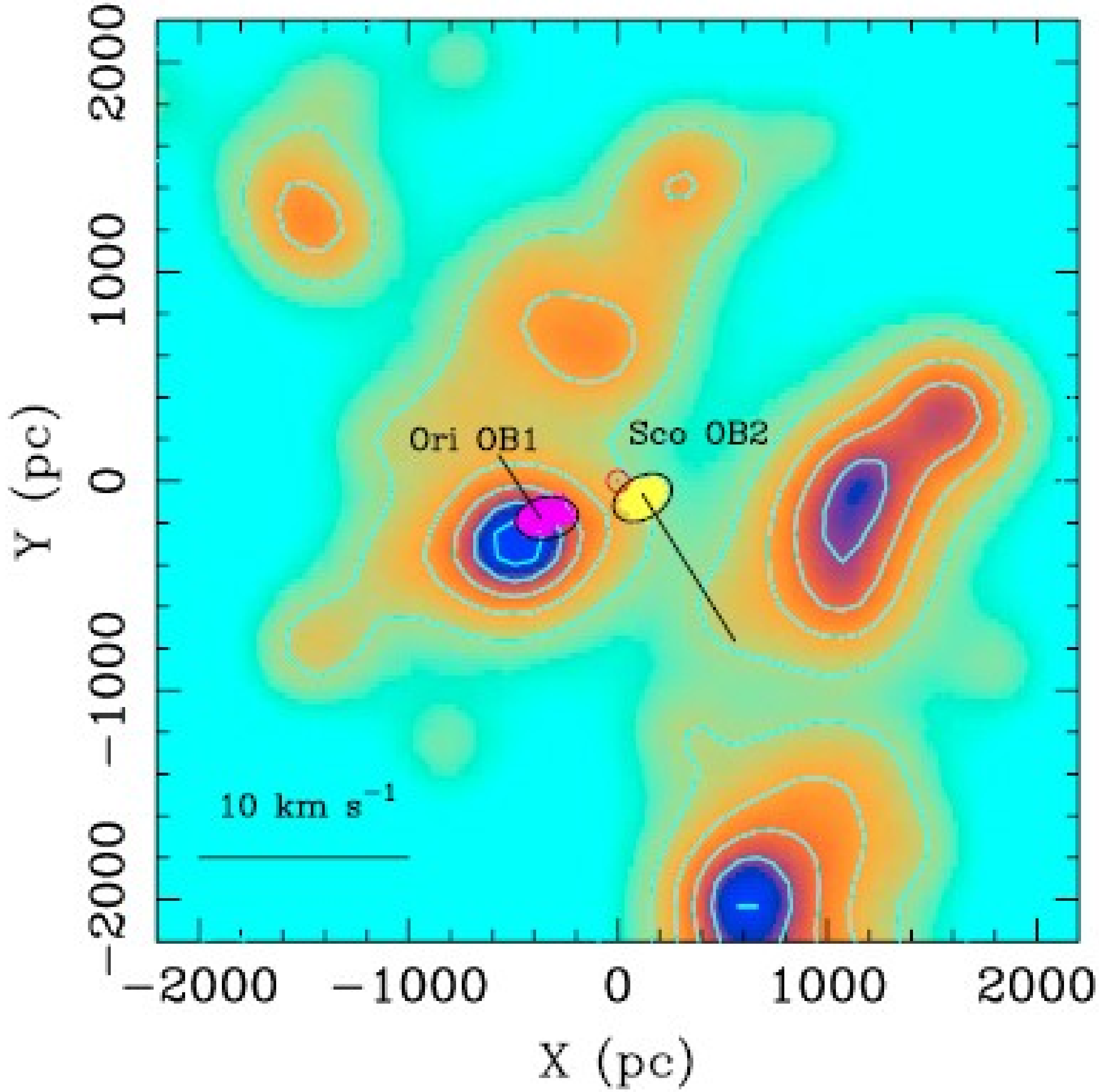
- We know of 14 Open Clusters which are younger than 10 Myrs within 1000 pc around the Sun (Source: WEBDA)
- There are also five star forming regions
- Open Clusters still have to form within the solar vicinity
- Total masses: up to 40 000 M(sun)
- Stable for some Gyrs
- Evolutionary theory has to explain these facts



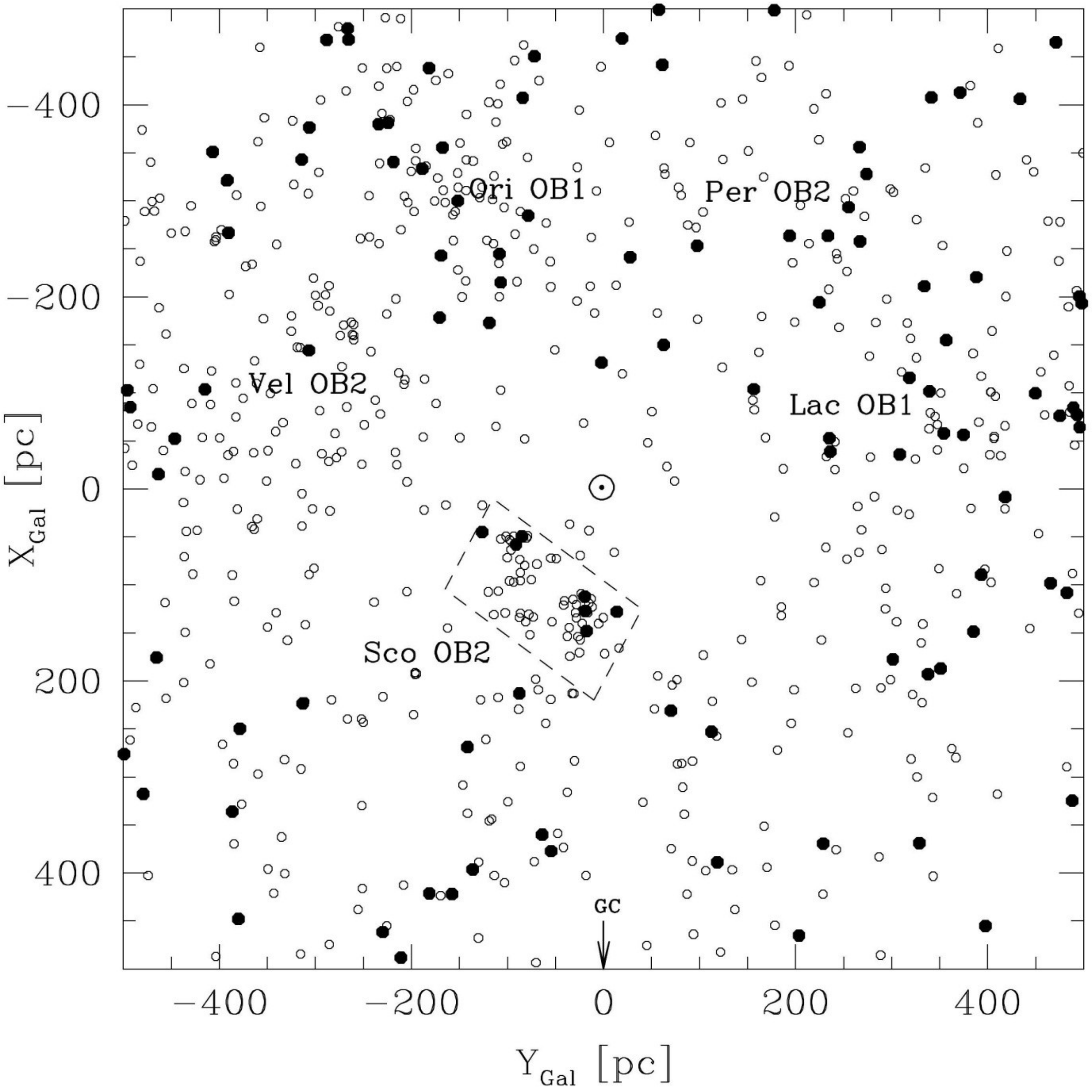
Clusters selected

Cluster_name	RA_2000_Dec		l	b	Dist	Mod	EB-V	Age	ST	Z	Diam	Fe/H	MRV	pm RA	pm Dec
Mamajek 1	08 42 06	-79 01 38	292.482	-21.654	97	4.93	0.00	6.9		-35.8	40.0		+16.1	-30.00	+27.80
Collinder 70	05 35 31	-01 06 00	205.03	-17.35	391	8.09	0.04	6.71		-116.6	180.0		19.49	0.36	-0.68
ASCC 24	06 28 44	-07 01 11	216.64	-8.23	400	8.44	0.14	6.96		-57.3	42.0		16.35	-5.55	-4.05
ASCC 16	05 24 35	+01 47 59	200.98	-18.35	460	8.59	0.09	6.93		-144.8	74.4		0.75	+0.75	-0.18
NGC 1980	05 35 24	-05 54 35	209.51	-19.60	550	8.86	0.05	6.67	B1	-184.5	25.2		25.34	0.83	-0.36
Bochum 14	18 02 00	-23 41 00	6.388	-0.499	578	13.48	1.508	6.996		-5.0	2.0				
NGC 2264	06 40 58	+09 53 42	202.936	2.196	667	9.28	0.051	6.954	O7	25.6	39.0	-0.15	+25.5	-1.13	-3.80
ASCC 122	22 33 14	+39 36 36	95.91	-15.90	700	9.53	0.10	6.98		-191.8	86.4		-8.17	-0.29	-4.19
Collinder 419	20 17 59	+40 43 12	78.07	2.79	740	10.40	0.34	6.85	B2	36.0	30.0		-8.19	-2.56	-6.99
ASCC 79	15 19 11	-60 43 47	320.04	-2.86	800	10.01	0.16	6.86		-39.9	62.4		4.03	-2.67	-4.10
IC 5146	21 53 24	+47 16 00	94.383	-5.495	852	11.49	0.593	6.00	B1	-81.6	20.0			-1.77	-1.70
Lynga 14	16 55 04	-45 14 00	340.919	-1.089	881	14.15	1.428	6.712		-16.7	3.0				
Ruprecht 119	16 28 15	-51 30 00	333.276	-1.879	956	11.67	0.570	6.853		-31.3	8.0			-1.15	-1.80
NGC 6383	17 34 48	-32 34 00	355.690	0.041	985	10.89	0.298	6.962	O7	0.7	20.0		+7.00	+1.58	-2.00

Several spurious entries like the “ASCC clusters”



Distribution of young open clusters and star forming regions from Alfaro et al., 2009, Ap&SS, 324, 141



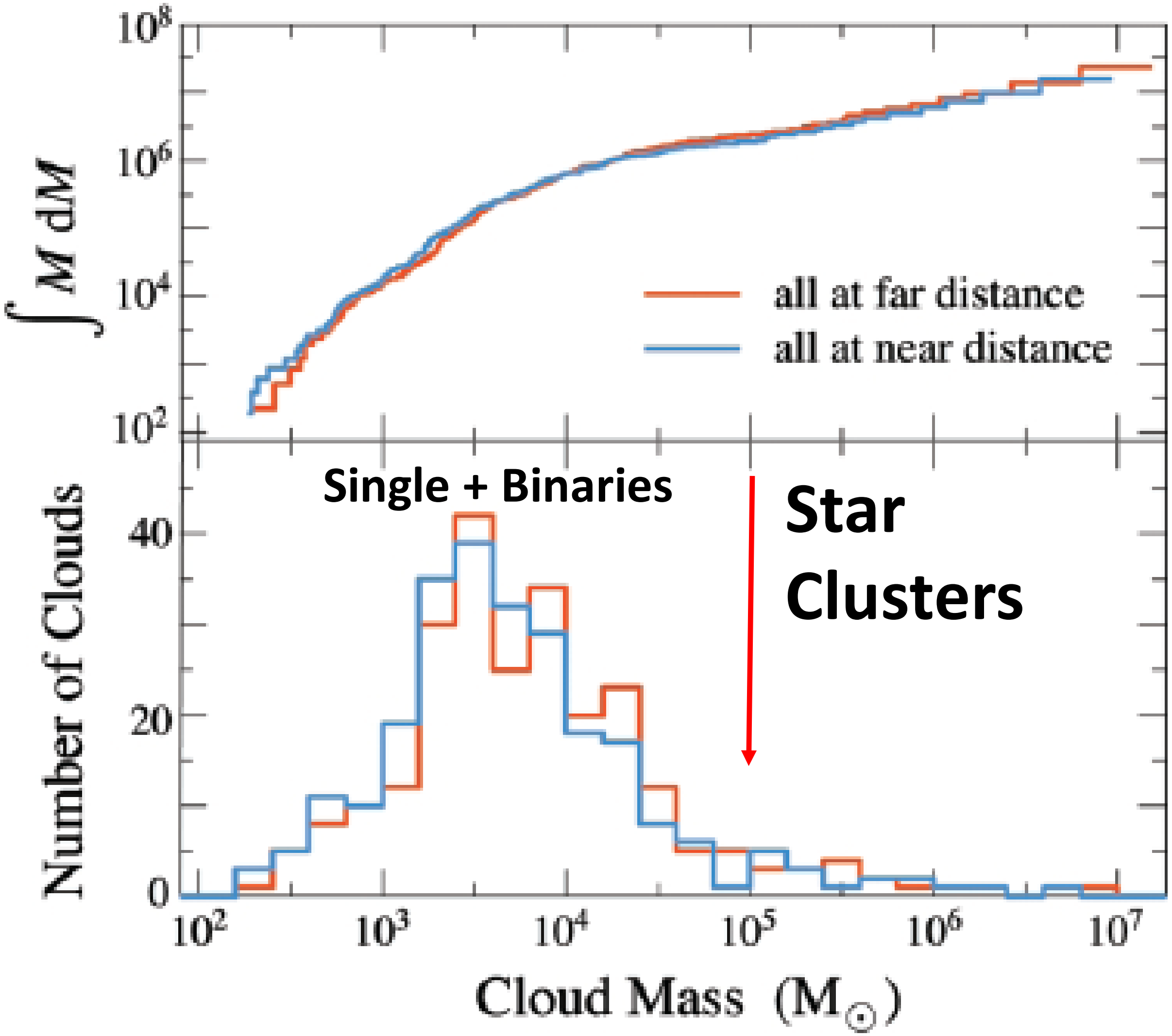
Stars hotter than B0 and B0 to B2

Distribution of star forming regions from Preibisch & Mamajek, 2008, Handbook of Star Forming Regions, Volume II

Giant Molecular Clouds

- Star Clusters can only form within „Giant Molecular Clouds“ (GMC) with a high enough initial mass
- The stellar formation rate in the solar neighborhood is very low
- But still there have to exist several GMCs to form Star Clusters
- Is the formation process the same for all observed Galaxy types?

Giant Molecular Clouds



Stark & Lee, 2006, ApJ, 641, L116

Single + Binaries

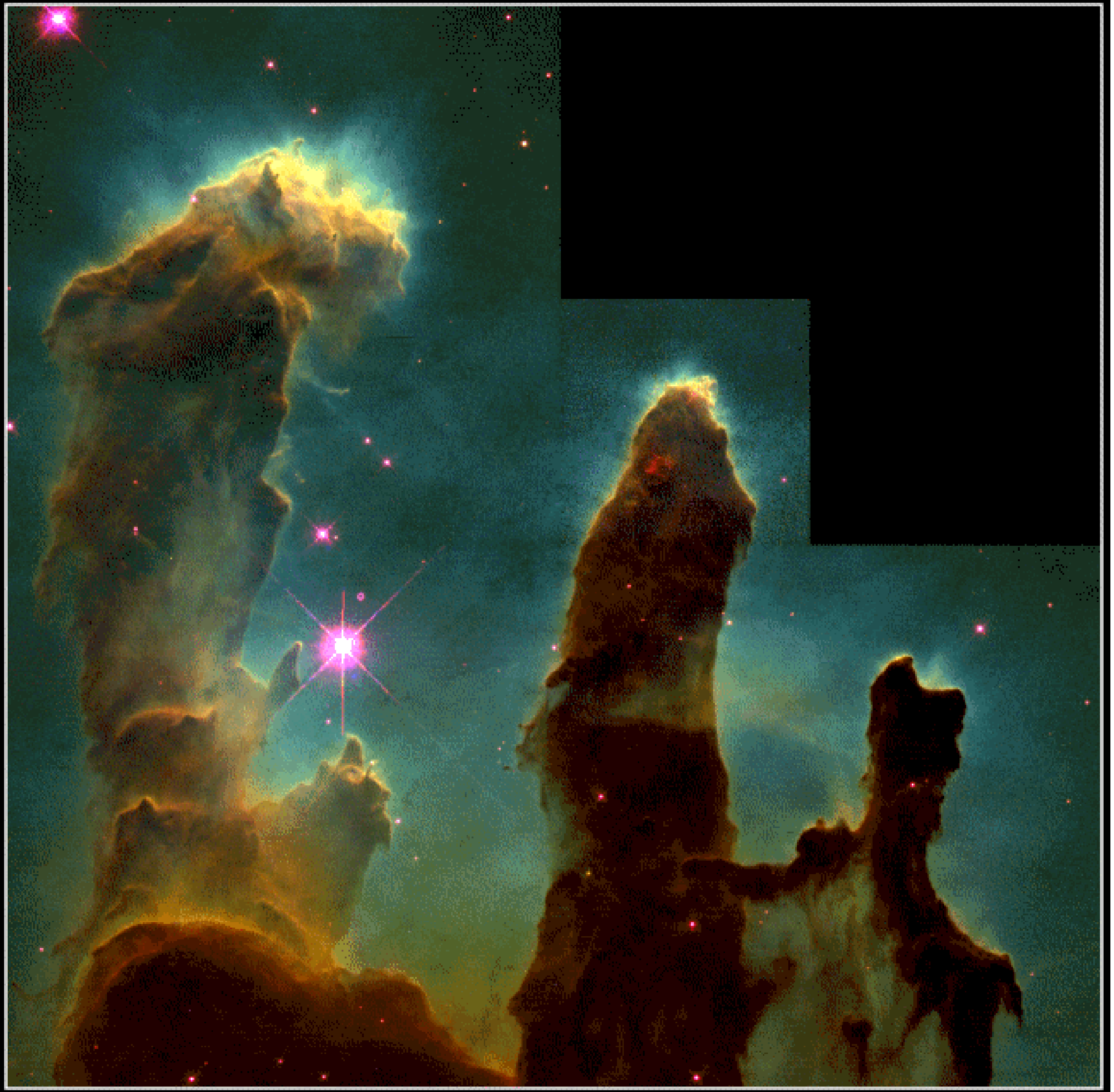
Star Clusters

Recent investigation of the ^{13}CO Gas within 2000 pc around the Sun

The number of young OCLs can be very well explained

Formation rate of 0.45 OCLs per kpc^{-2} Myr^{-1} in the galactic disk within 2 kpc around the Sun

Battinelli & Capuzzo-Dolcetta, 1991, MNRAS, 248, 76



Gaseous Pillars · M16 HST · WFPC2

PRC95-44a · ST ScI OPO · November 2, 1995
J. Hester and P. Scowen (AZ State Univ.), NASA

NGC 6611 (M16)

d = 1750 pc
t = 8 Myr

Star formation „live“

Initial Mass Function

- The „Initial Mass Function“ (IMF) describes the mass distribution for a population of stars when they are formed together
- Relevant astrophysics:
  1. Size, total mass and metallicity of the initial GMC
  2. Fragmentation of the GMC
  3. Conservation of the angular momentum
  4. Local and global magnetic fields
  5. Accretion in the Pre-Main Sequence phase
- The only observational parameter for the test of stellar formation and evolution models
- We observe a luminosity function which has to be transformed to the IMF

Initial Mass Function

- Several most important questions are still not solved
  1. Is the IMF homogeneous within the Milky Way?
  2. Is the IMF constant throughout time?
  3. What is the influence of the local and global magnetic field on the IMF?
  4. What is the influence of the local and global metallicity on the IMF?

Initial Mass Function

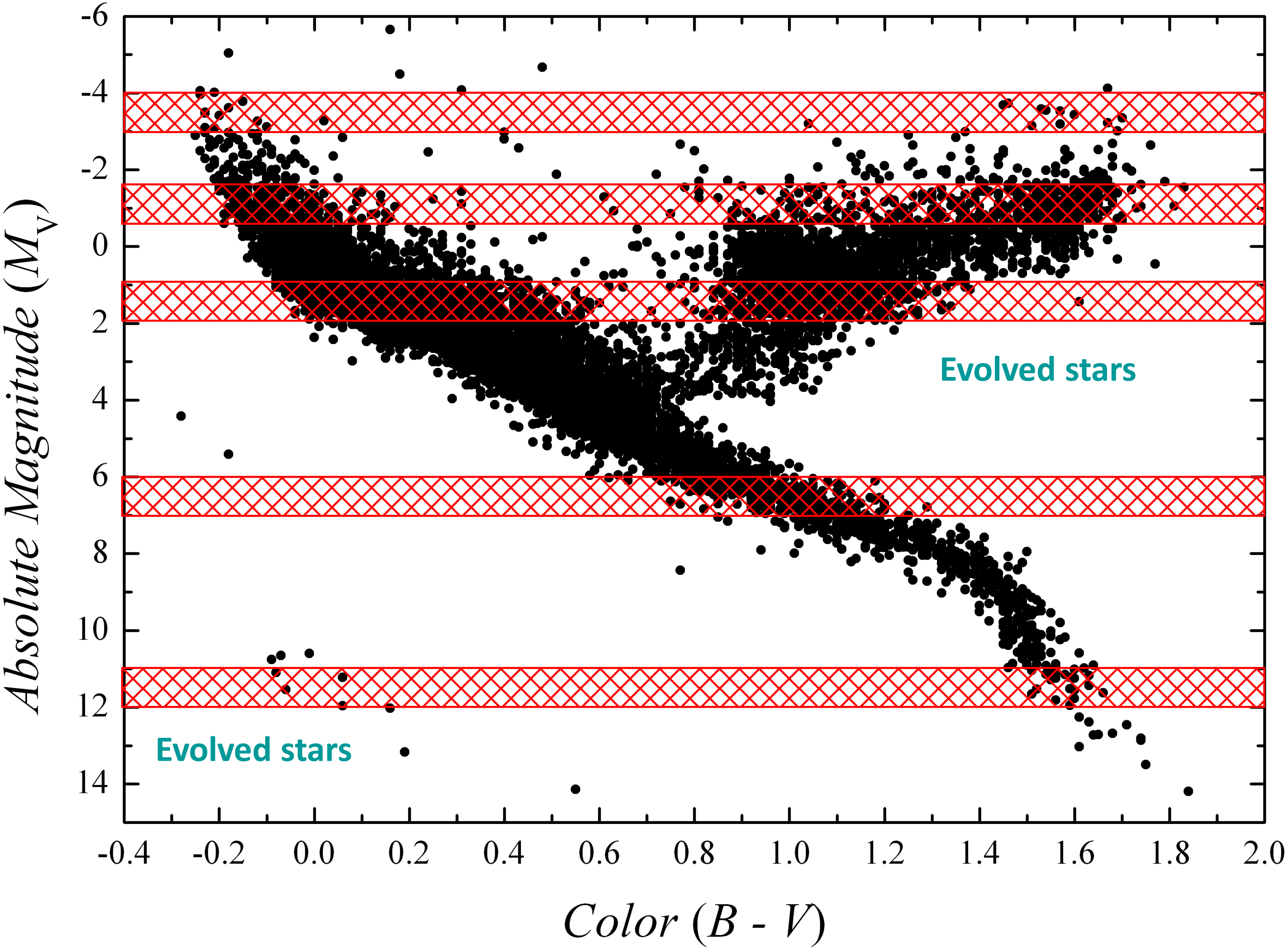
The IMF $\theta(m)$, often called „Present-Day Mass Function“ (PDMF), is defined as:

$$dN = \theta(m)\,dm$$

dN is the number of all stars per cubic parsec on the *main sequence* with a mass between M and (M + dm).

But we observe not the masses of stars but their magnitudes (relative and absolute) or luminosities.

So we have to define the luminosity function and transform it into the IMF.



In each row (M_V + dM) there is a mixture of main sequence and evolved objects. For the IMF, we need the main sequence only.

Luminosity function

The luminosity function $\Psi(M_V)$, is defined as:

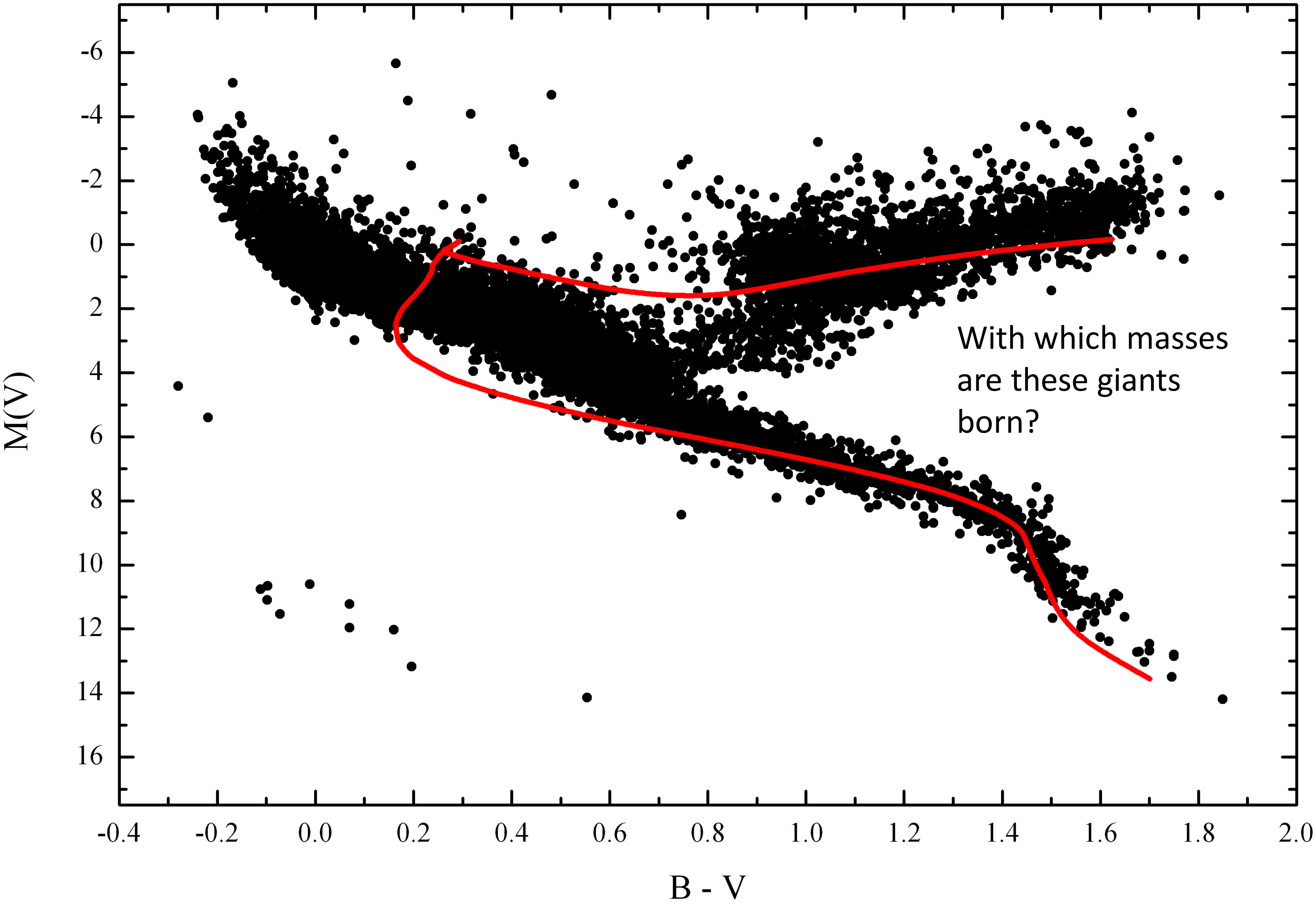
$$dN = -\Psi(M_V)\, dM_V$$

dN is is the number of all stars per cubic parsec on the *main sequence* with an absolute magnitude between M_V and $(M_V + dM_V)$. The transformation to the IMF is given as:

$$\theta(m) = -\Psi(M_V)[dm(M_V)/dM_V]^{-1}$$

The second term is the derivation of the Mass-Luminosity function $m(M_V)$. It is depending on the age (t), metallicity (Z) and rotation (v_{rot})

$$m(M_V) = m(M_V, Z, t, v_{rot})$$



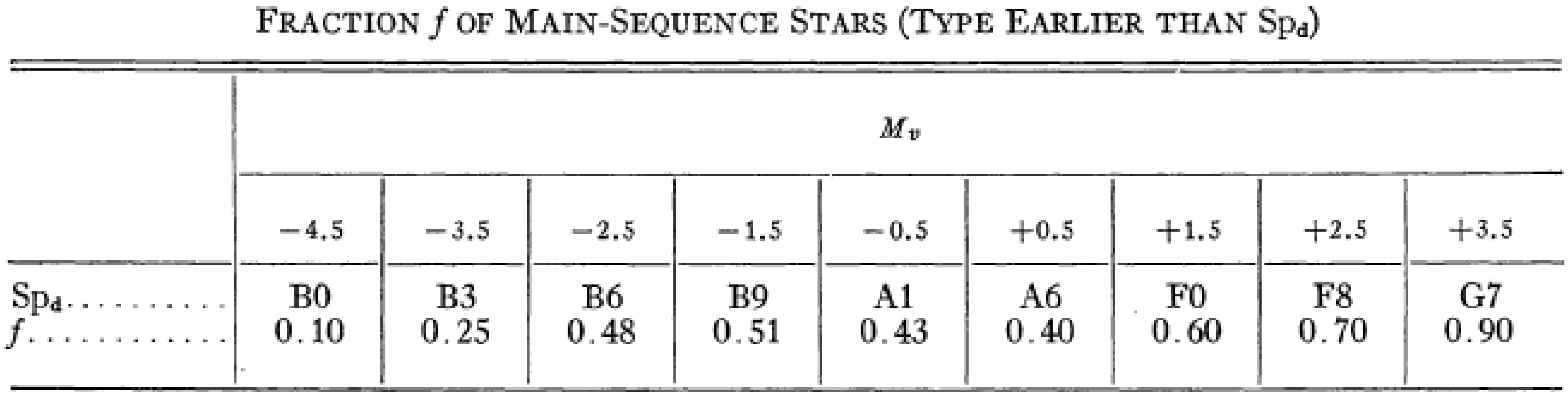
With which masses are these giants born?

Correction of the observations

We have to correct the complete observations for the evolved objects. There are three possibilities:

1. Take a statistical sample with a well known luminosity function (clusters)
2. Take a statistical sample with well known photometric magnitudes and distances
3. Take isochrones = theoretical star evolution = models based on observations = circular argument

All these methods are not self consistent and always introduce an unknown error to the analysis



FRACTION f OF MAIN-SEQUENCE STARS (TYPE EARLIER THAN Sp_d)

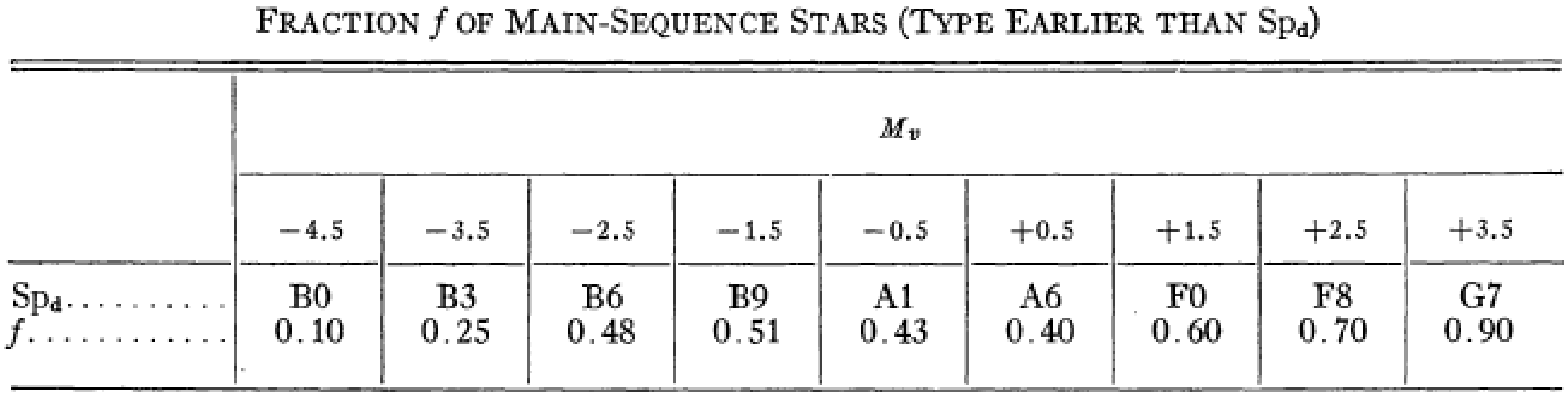
	M_v								
	−4.5	−3.5	−2.5	−1.5	−0.5	+0.5	+1.5	+2.5	+3.5
Sp_d	B0	B3	B6	B9	A1	A6	F0	F8	G7
f	0.10	0.25	0.48	0.51	0.43	0.40	0.60	0.70	0.90

Salpeter, 1955, ApJ, 121, 161

Results of classical spectral classification, only 10% of stars with M_V = -4.5 mag are on the main sequence!

These values are depending on the chosen sample for the spectral classification and which classification scheme is applied.

The errors are rather large.



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Results of classical spectral classification, only 10% of stars with M_V = -4.5 mag are on the main sequence!

These values are depending on the chosen sample for the spectral classification and which classification scheme is applied.

The errors are rather large.

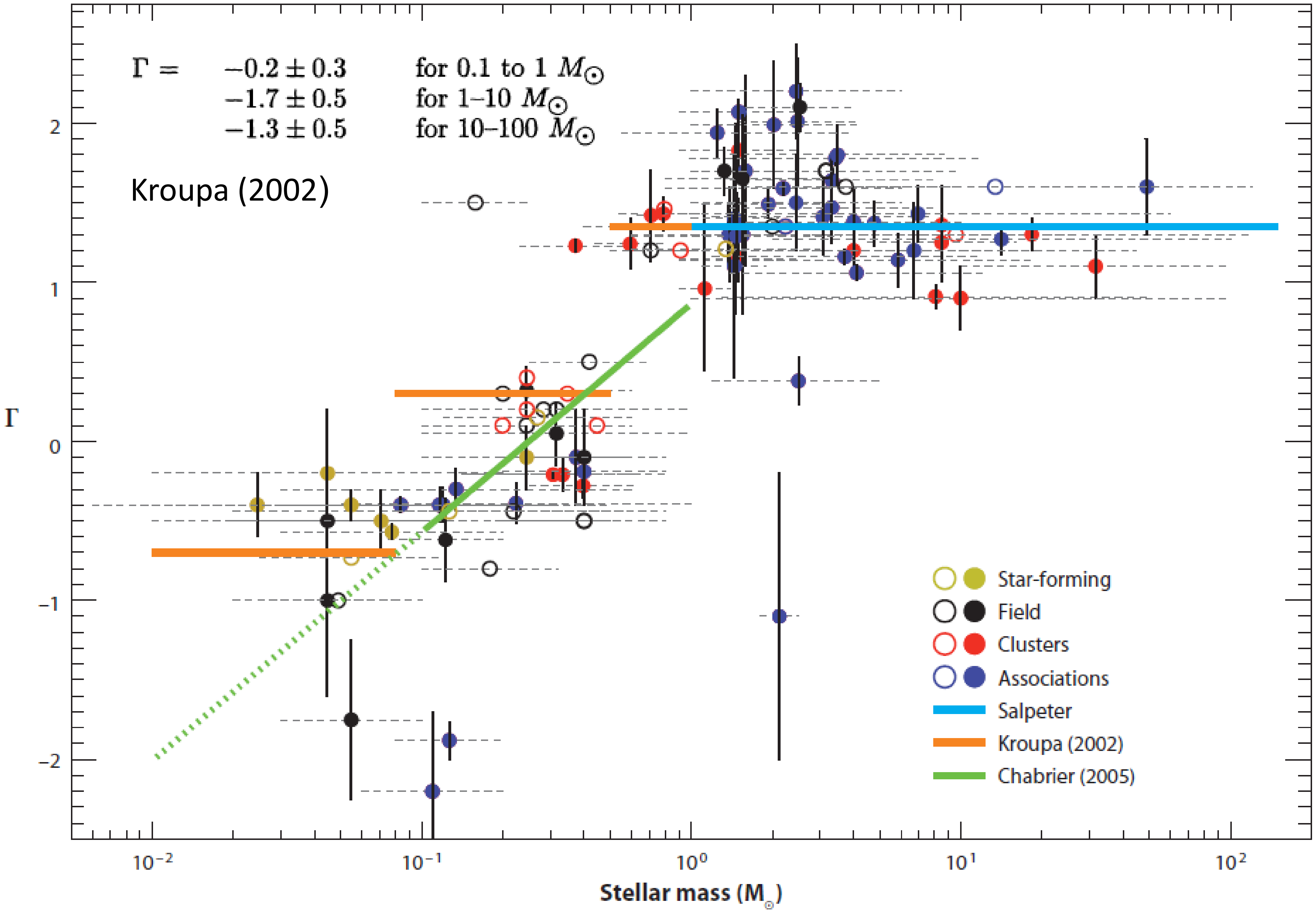
All observations have to be normalized to one “standard system” which means essentially to one “time scale”.

The observations show, that this heuristic law describes them very well

$\theta(m) \approx m^{-\Gamma}$ Salpeter law (1955)

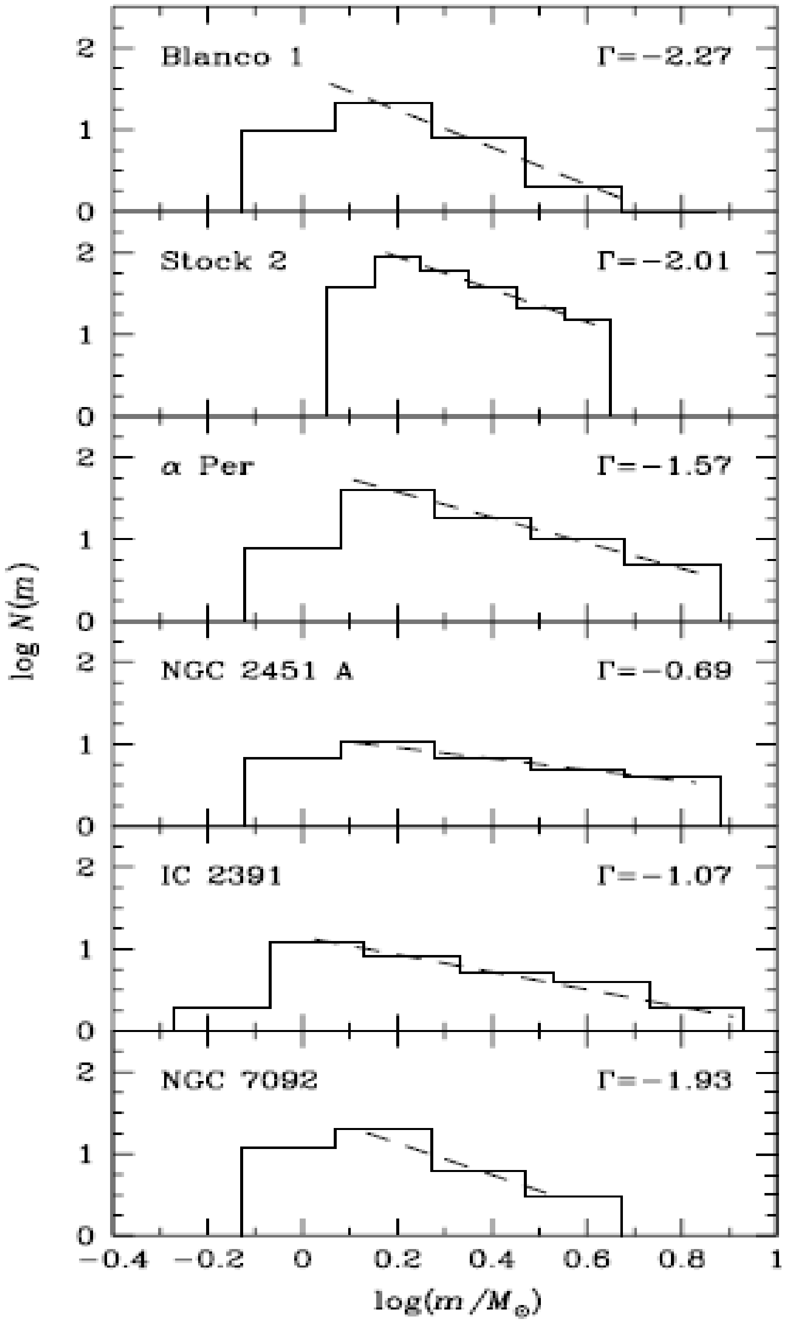
Star cluster are one of the most important observational test for the IMF because they, normally, have well defined ages, distances and metallicities. However, the errors are still quiet large.

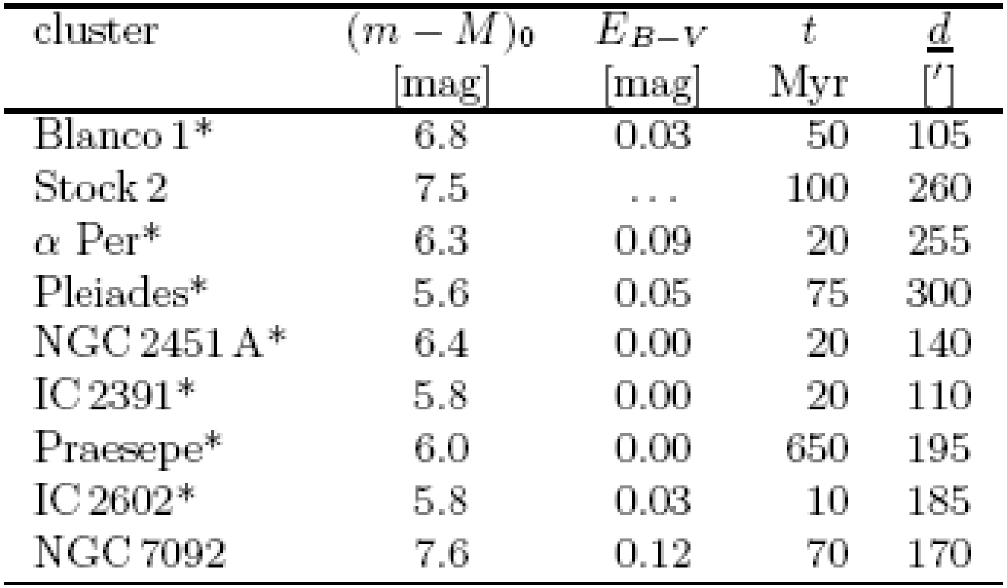
But there is still no homogeneous IMF determination for open clusters taking into account the available data.



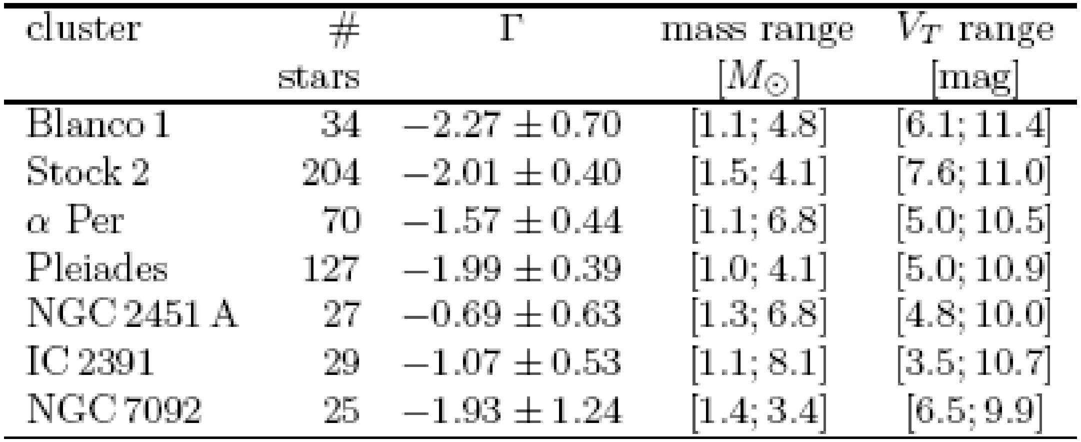
Bastian et al, 2010, Annual Review of Astronomy and Astrophysics, 48, 339

TYCHO2 data





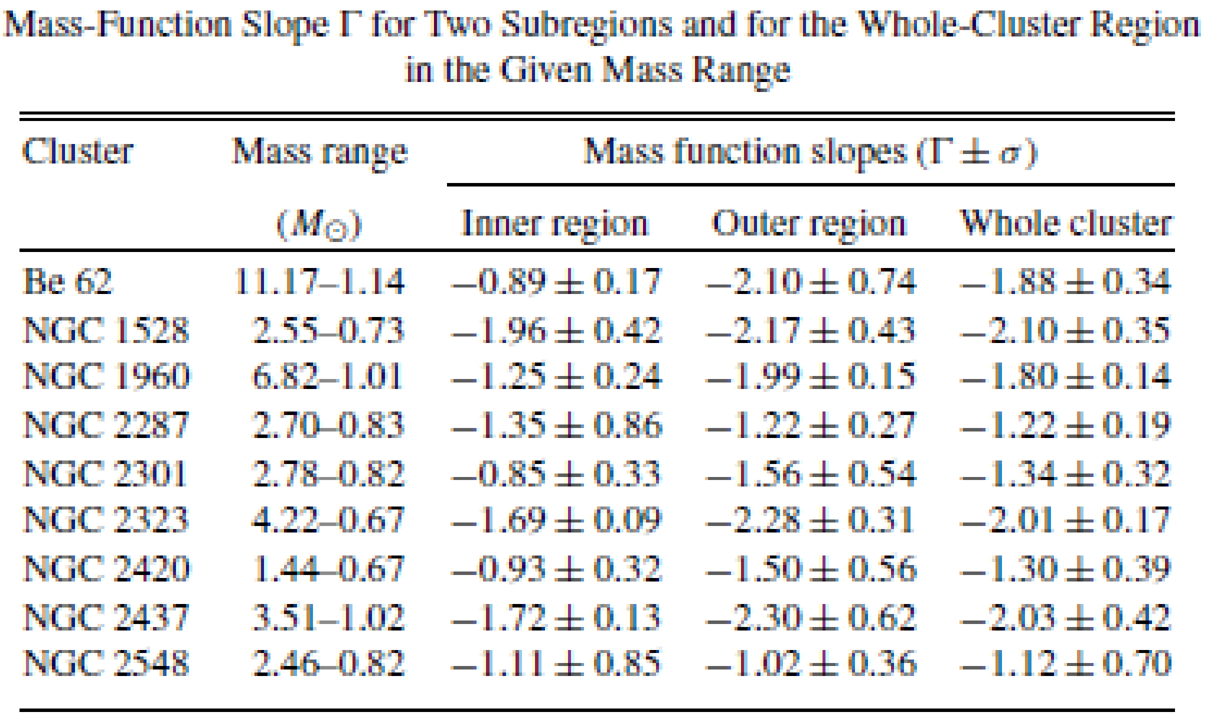
cluster	$(m-M)_0$ [mag]	E_{B-V} [mag]	t Myr	$\underline{d}$ [′]
Blanco 1*	6.8	0.03	50	105
Stock 2	7.5	…	100	260
α Per*	6.3	0.09	20	255
Pleiades*	5.6	0.05	75	300
NGC 2451 A*	6.4	0.00	20	140
IC 2391*	5.8	0.00	20	110
Praesepe*	6.0	0.00	650	195
IC 2602*	5.8	0.03	10	185
NGC 7092	7.6	0.12	70	170



cluster	# stars	Γ	mass range [$M_\odot$]	V_T range [mag]
Blanco 1	34	-2.27 ± 0.70	[1.1; 4.8]	[6.1; 11.4]
Stock 2	204	-2.01 ± 0.40	[1.5; 4.1]	[7.6; 11.0]
α Per	70	-1.57 ± 0.44	[1.1; 6.8]	[5.0; 10.5]
Pleiades	127	-1.99 ± 0.39	[1.0; 4.1]	[5.0; 10.9]
NGC 2451 A	27	-0.69 ± 0.63	[1.3; 6.8]	[4.8; 10.0]
IC 2391	29	-1.07 ± 0.53	[1.1; 8.1]	[3.5; 10.7]
NGC 7092	25	-1.93 ± 1.24	[1.4; 3.4]	[6.5; 9.9]

Sanner & Geffert, 2001, A&A, 370, 87

Saurabh et al., 2008, AJ, 135, 1934



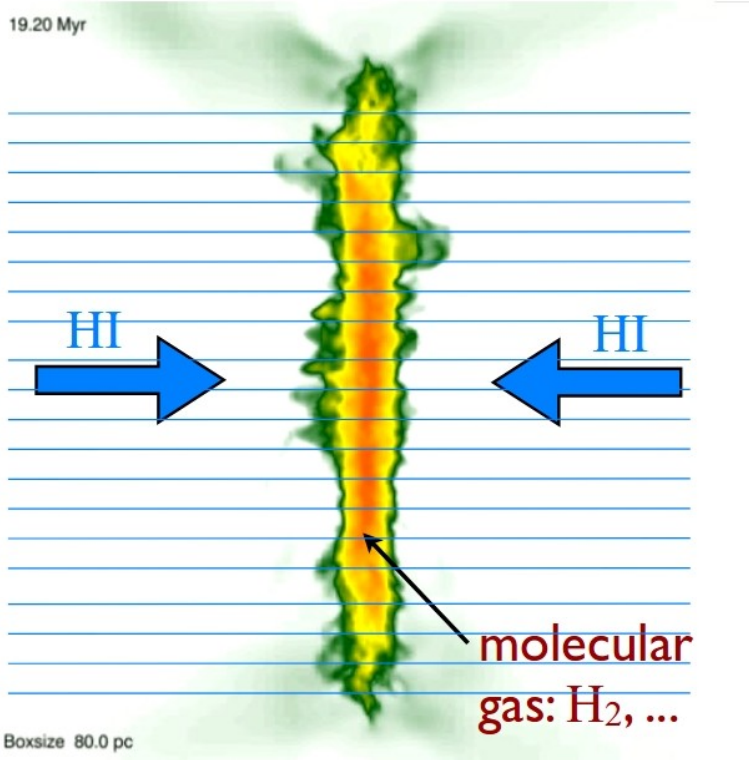
Mass-Function Slope Γ for Two Subregions and for the Whole-Cluster Region in the Given Mass Range

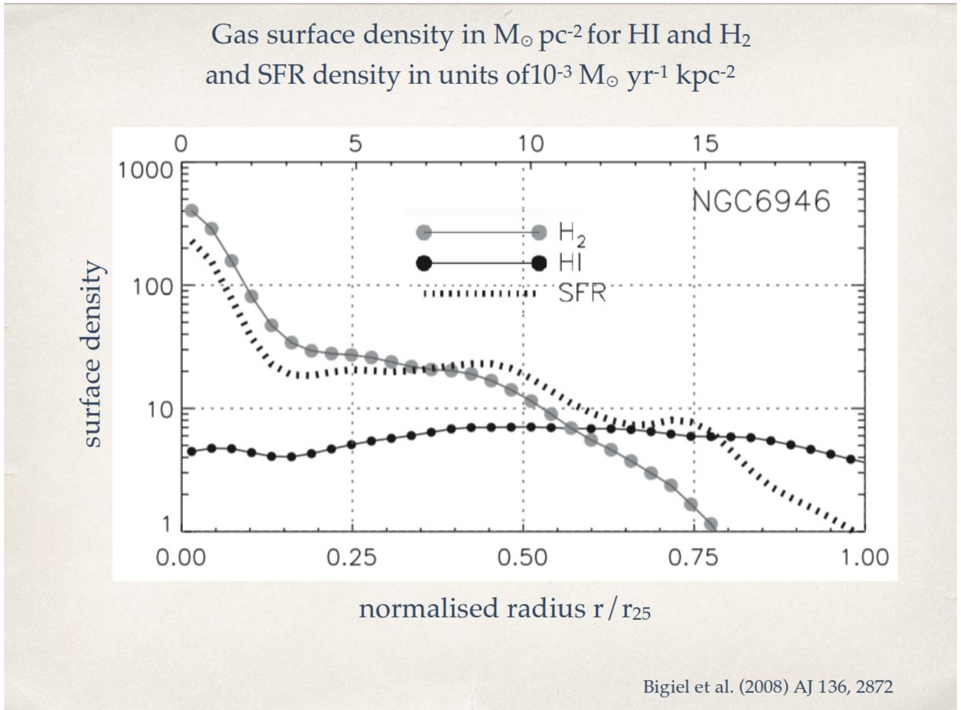
Cluster	Mass range ($M_\odot$)	Mass function slopes ($\Gamma \pm \sigma$) Inner region	Outer region	Whole cluster
Be 62	11.17–1.14	-0.89 ± 0.17	-2.10 ± 0.74	-1.88 ± 0.34
NGC 1528	2.55–0.73	-1.96 ± 0.42	-2.17 ± 0.43	-2.10 ± 0.35
NGC 1960	6.82–1.01	-1.25 ± 0.24	-1.99 ± 0.15	-1.80 ± 0.14
NGC 2287	2.70–0.83	-1.35 ± 0.86	-1.22 ± 0.27	-1.22 ± 0.19
NGC 2301	2.78–0.82	-0.85 ± 0.33	-1.56 ± 0.54	-1.34 ± 0.32
NGC 2323	4.22–0.67	-1.69 ± 0.09	-2.28 ± 0.31	-2.01 ± 0.17
NGC 2420	1.44–0.67	-0.93 ± 0.32	-1.50 ± 0.56	-1.30 ± 0.39
NGC 2437	3.51–1.02	-1.72 ± 0.13	-2.30 ± 0.62	-2.03 ± 0.42
NGC 2548	2.46–0.82	-1.11 ± 0.85	-1.02 ± 0.36	-1.12 ± 0.70

Typical values and errors

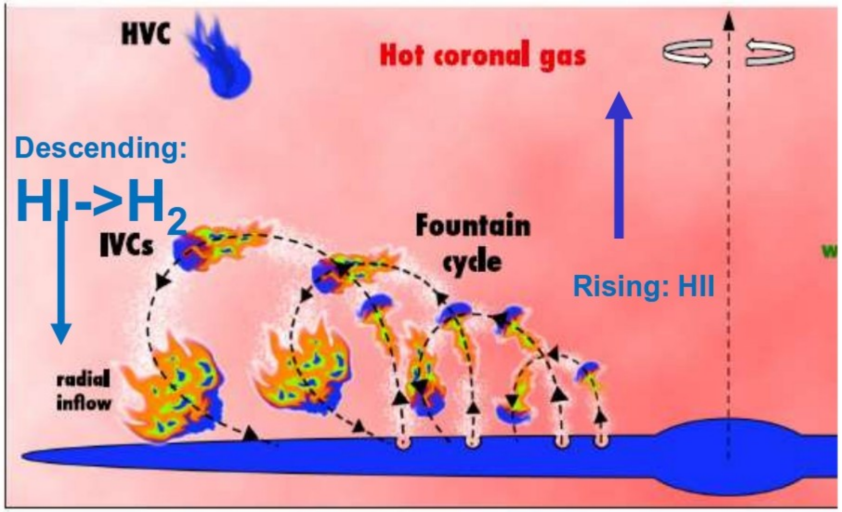
Molecular clouds - formation

- Colliding flows – colliding HI streams (only low-mass cloud formation – 10^4-10^5 M_{sol})
- Cloud collisions – in spiral arms (can yield rather massive clouds - 10^6 M_{sol})
- Various instabilities ? - can yield massive clouds
- Subsequent supernova explosions – sweeping up the local ISM
- **In all cases, HII formation is needed**




Gas surface density in M⊙ pc⁻² for HI and H₂ and SFR density in units of10⁻³ M⊙ yr⁻¹ kpc⁻²

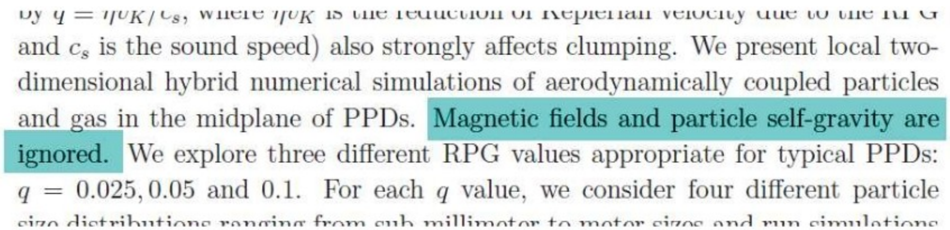
Bigiel et al. (2008) AJ 136, 2872



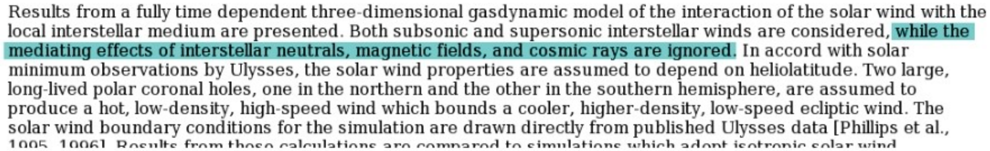
Marasco, Marinacci & Fraternali 2013, MNRAS 443, 1634

Magnetic fields

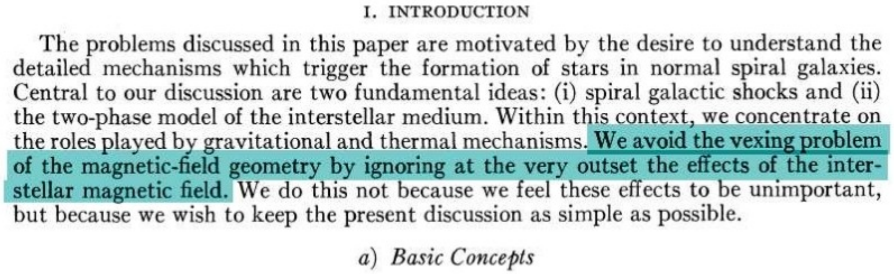
Slide credit: Kothes (2018)



by $q = \eta v_K / c_s$, where ηv_K is the reduction of Keplerian velocity due to the RPG and c_s is the sound speed) also strongly affects clumping. We present local two-dimensional hybrid numerical simulations of aerodynamically coupled particles and gas in the midplane of PPDs. Magnetic fields and particle self-gravity are ignored. We explore three different RPG values appropriate for typical PPDs: $q = 0.025, 0.05$ and 0.1. For each q value, we consider four different particle size distributions ranging from sub-millimeter to meter sizes and run simulations



Results from a fully time dependent three-dimensional gasdynamic model of the interaction of the solar wind with the local interstellar medium are presented. Both subsonic and supersonic interstellar winds are considered, while the mediating effects of interstellar neutrals, magnetic fields, and cosmic rays are ignored. In accord with solar minimum observations by Ulysses, the solar wind properties are assumed to depend on heliolatitude. Two large, long-lived polar coronal holes, one in the northern and the other in the southern hemisphere, are assumed to produce a hot, low-density, high-speed wind which bounds a cooler, higher-density, low-speed ecliptic wind. The solar wind boundary conditions for the simulation are drawn directly from published Ulysses data [Phillips et al., 1995, 1996]. Results from these calculations are compared to simulations which adopt isotropic solar wind



I. INTRODUCTION

The problems discussed in this paper are motivated by the desire to understand the detailed mechanisms which trigger the formation of stars in normal spiral galaxies. Central to our discussion are two fundamental ideas: (i) spiral galactic shocks and (ii) the two-phase model of the interstellar medium. Within this context, we concentrate on the roles played by gravitational and thermal mechanisms. We avoid the vexing problem of the magnetic-field geometry by ignoring at the very outset the effects of the interstellar magnetic field. We do this not because we feel these effects to be unimportant, but because we wish to keep the present discussion as simple as possible.

a) Basic Concepts

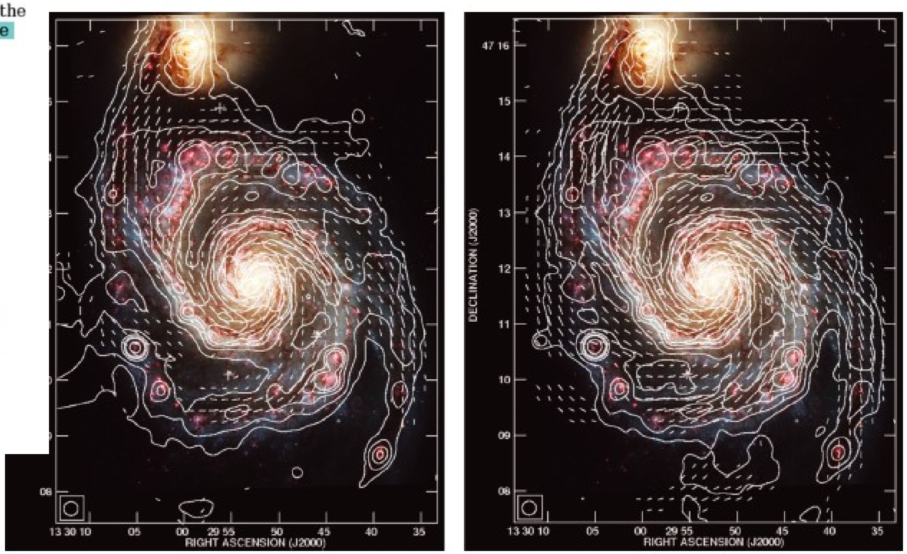
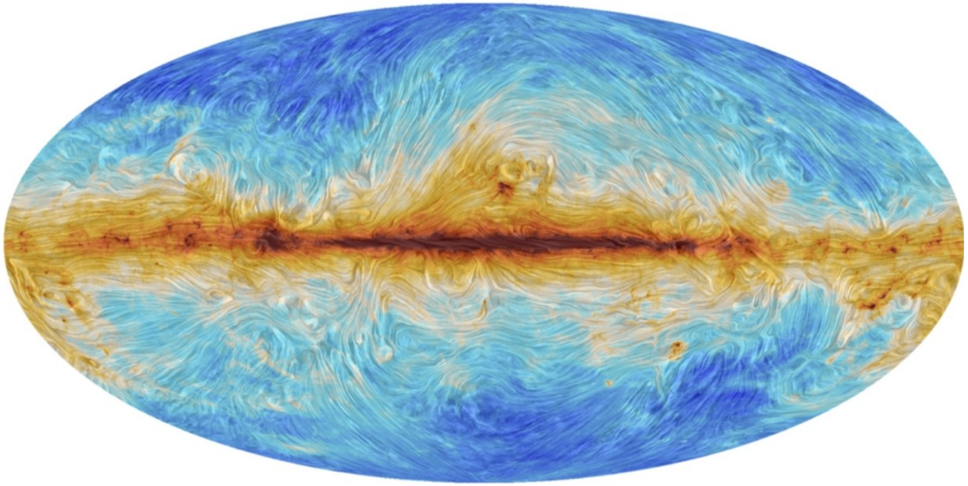


Figure 1. (a) λ3 cm (left) and (b) λ6 cm (right) radio emission at 15 arcsec resolution from VLA and Effelsberg observations, overlaid on a *Hubble Space Telescope* optical image [image credit: NASA, ESA, S. Beckwith (STScI) and The Hubble Heritage Team (STScI/AURA)]. Total intensity contours in both maps are at 6, 12, 24, 36, 48, 96, 192 times the noise levels of 20 μJy beam^{-1} at λ3 cm and 30 μJy beam^{-1} at λ6 cm. (Note that the roughly horizontal contours at the left edge of panel (a) are artefacts arising from mosaicking the two VLA pointings.) Also shown are the *B*-vectors of polarized emission: the plane of polarization of the observed electric field rotated by 90°, not corrected for Faraday rotation, with a length proportional to the polarized intensity (PI) and only plotted where PI $\geq 3\sigma_{\rm PI}$.

M51, Fletcher+ (2011)

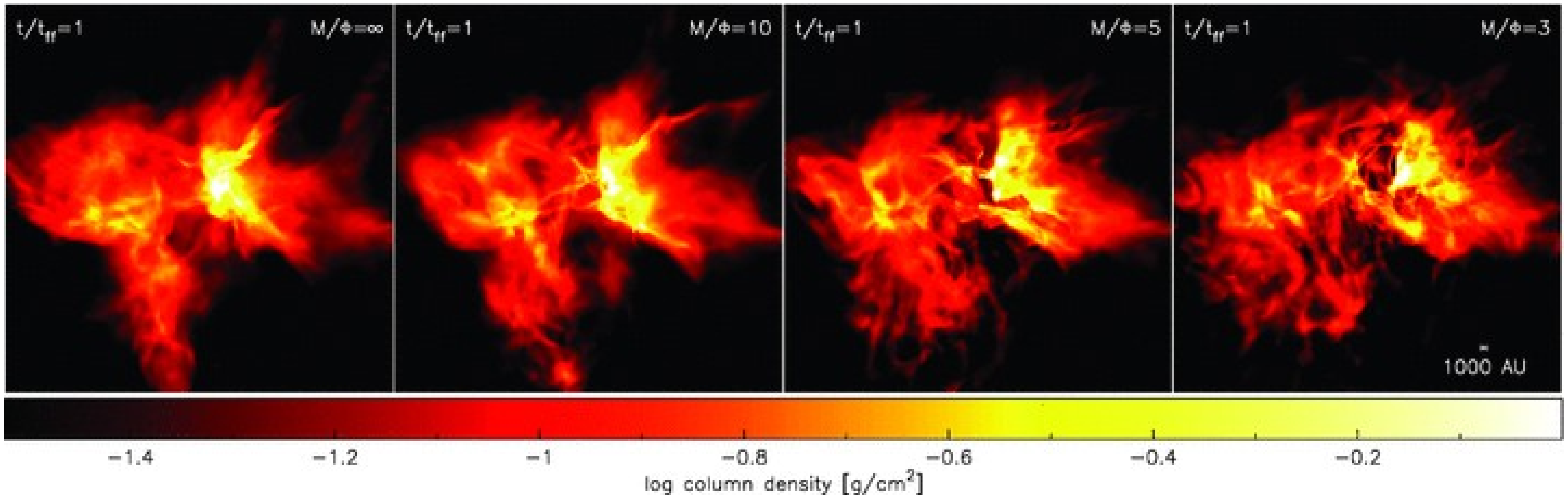


ESA PLANCK: *Milky Way's magnetic fingerprint (2015)*

Magnetic field of the Milky way from dust polarization

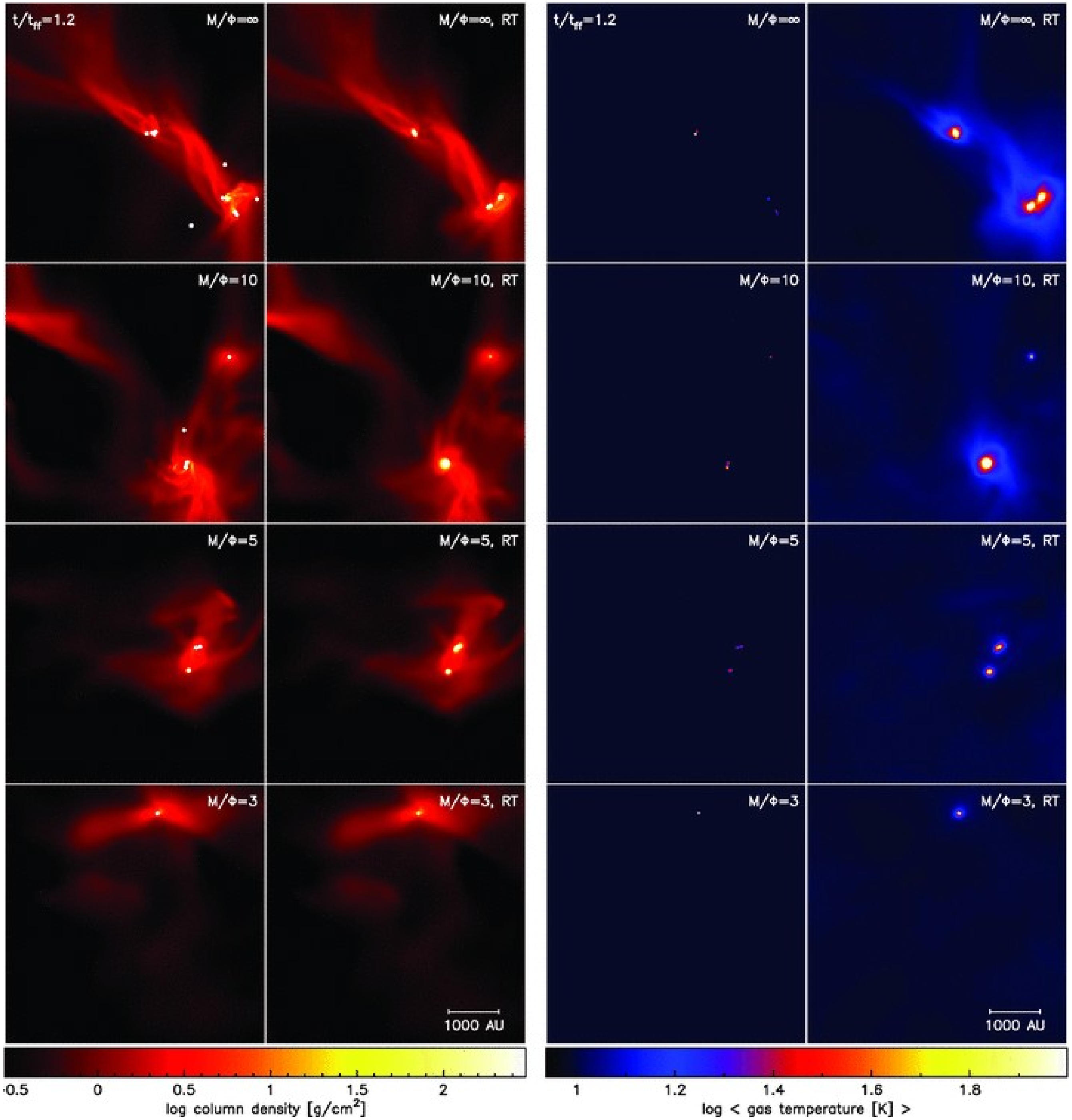
Magnetic field – star formation

- Price & Bate, 2009, MNRAS, 398, 33
- Effects of magnetic pressure on fragmentation



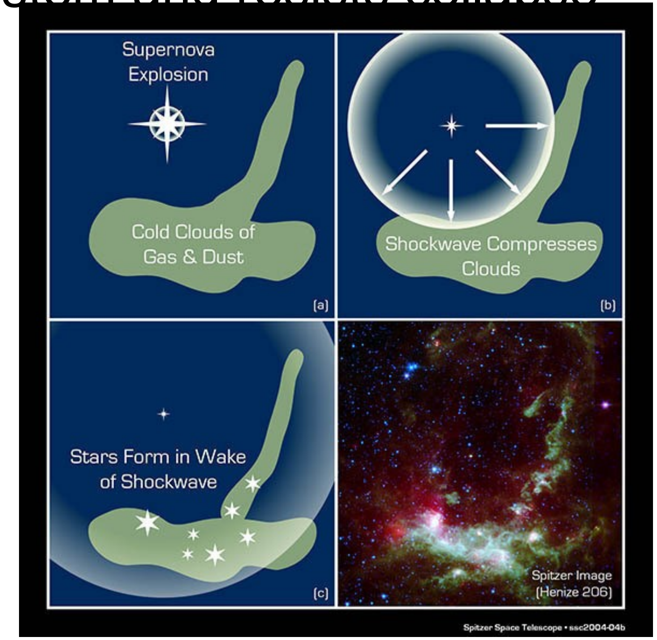
Increasing magnetic field strength

Increasing magnetic field strength



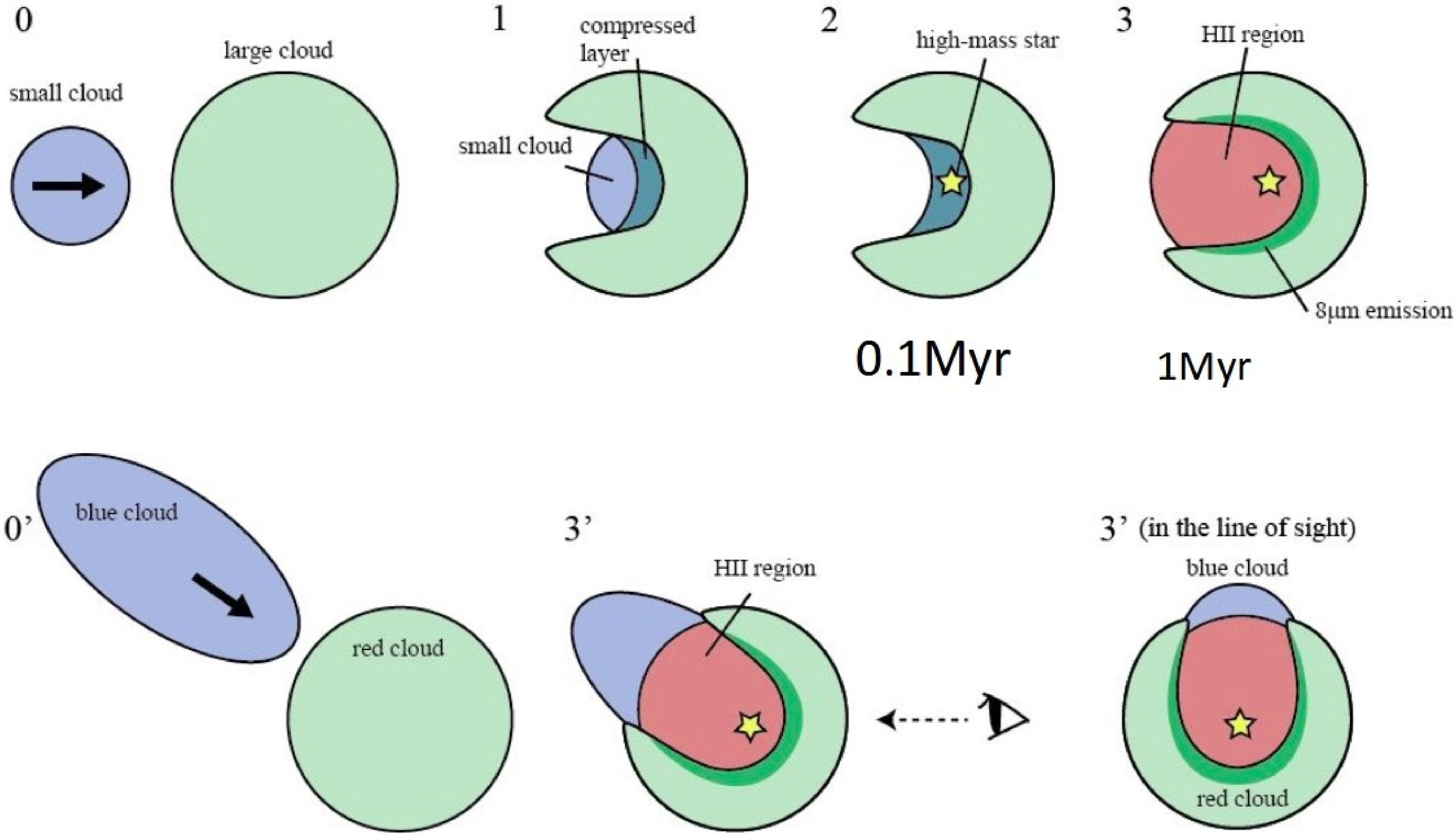
Star formation physics

- Inertia – turbulence effects
- Inertia – centrifugal force
- Fission – break-up of the collapsing cloud
- Heat pressure – warm gas tends to expand under its own pressure
- Magnetic pressure – mag. field amplified when compressed, changes the dynamics of the system and resists collapse



Constituent	Roles	Produced by	Destroyed by
Dense gas clouds	**Form stars**	**Gravitational collapse; supernova explosions?**	**Ultraviolet starlight; stellar winds, supernovae**
Dust grains	**Catalyze molecule formation; stop ultraviolet starlight**	**Red giant stars**	**Supernovae**
Molecules	**Radiate heat from gas clouds, permitting collapse**	**Grain catalysis**	**Ultraviolet starlight**
Stars	**Produce supernovae, red giants, dust, ultraviolet light**	**Gravitational collapse**	**Stellar evolution**

Cloud-cloud collision



0.1Myr

1Myr

Credit: Fukui (2018)

Collect & collapse

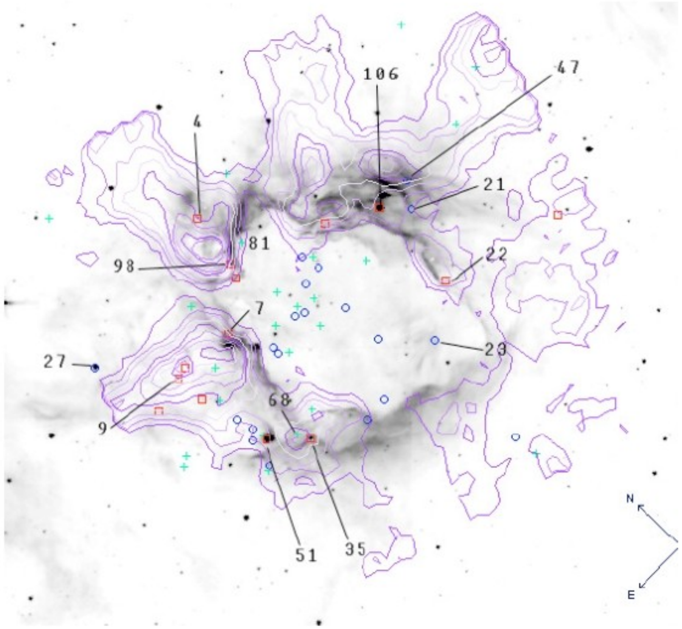


Figure 4: Example of a molecular shell surrounding an HII region with presence of YSOs. The image was extracted from Pomarès et al. (2009). The borders of the HII region are traced with the 8 μm emission (in gray) and the molecular shell, shown in contours, is mapped through the emission of ^{13}CO. Red squares are the positions of Class I sources, i.e. the youngest YSOs.

During a supersonic expansion of HII region, enough of ISM can be swept up to initialize SF

Radiatively-driven implosion

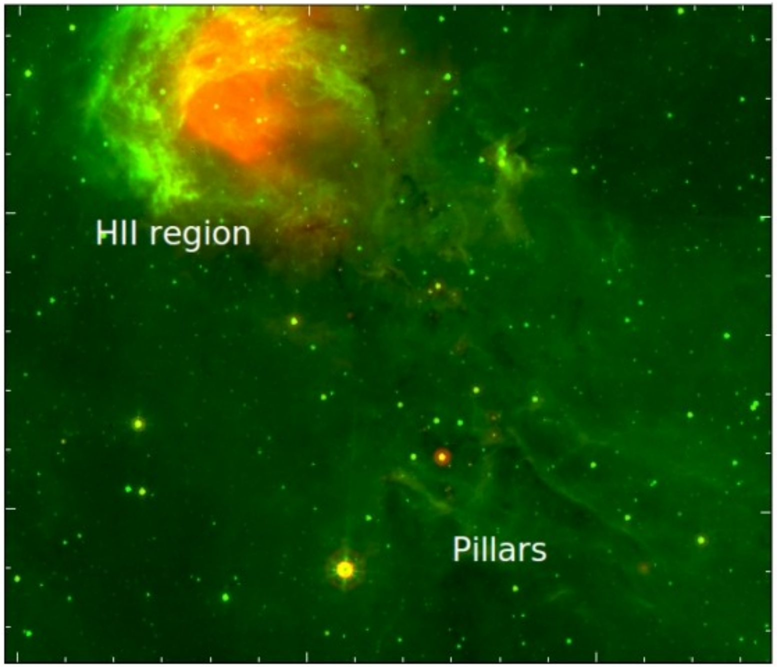
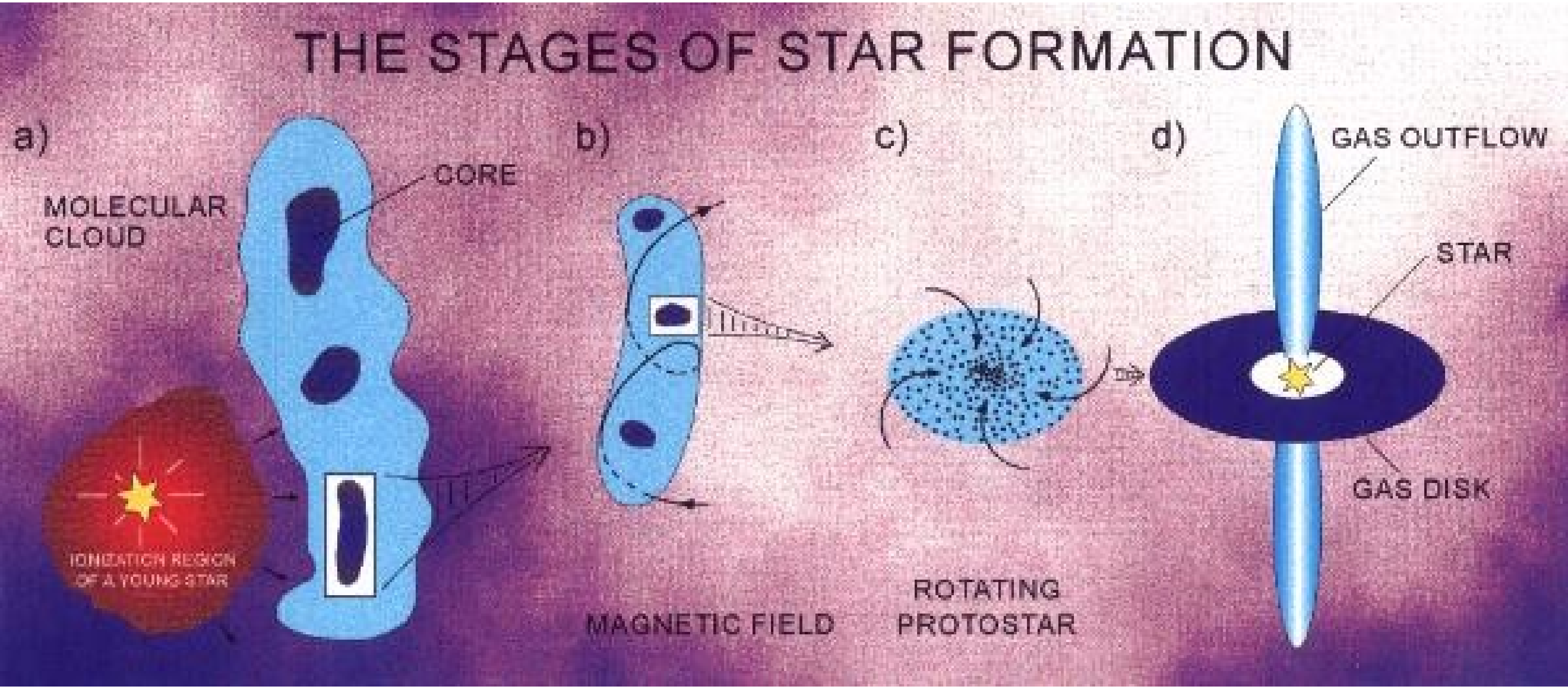


Figure 5: Example of an HII region irradiating pillar-like shape molecular condensations. The 8 and 24 μm emissions are displayed in green and red, respectively. From Paron et al. (2017).

Ionization from from a HII region moves over a molecular condensation and generates a dense outer shell of ionized gas. This shell is over-pressured with respect to the interior of the condenstation and shocks are driven into it, compressing the interior until the pressure is balanced

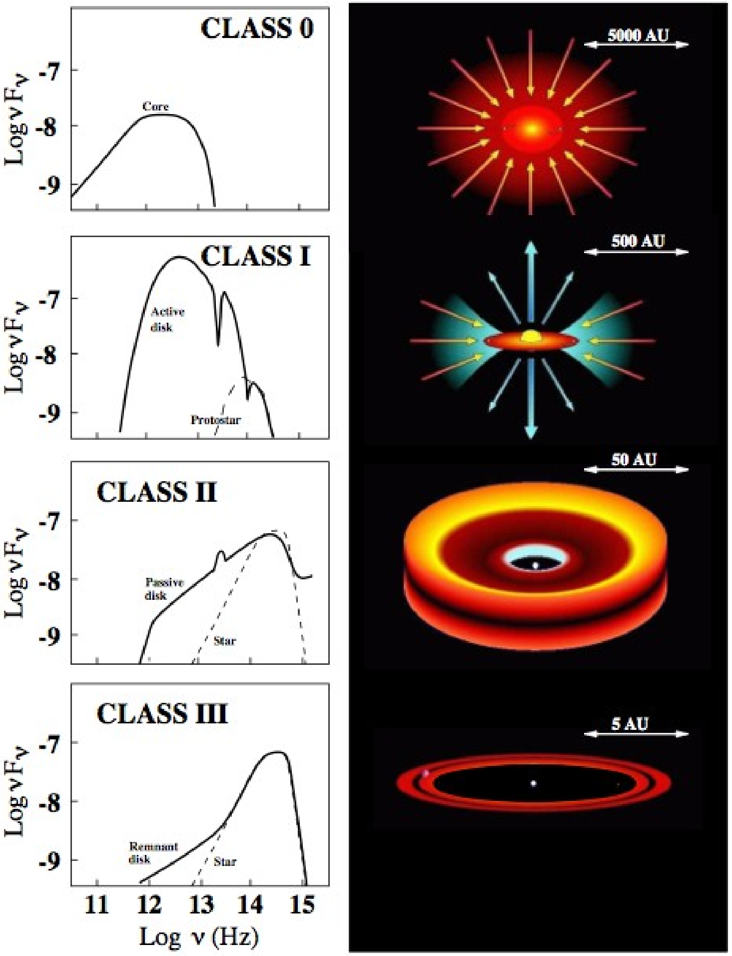
Star formation



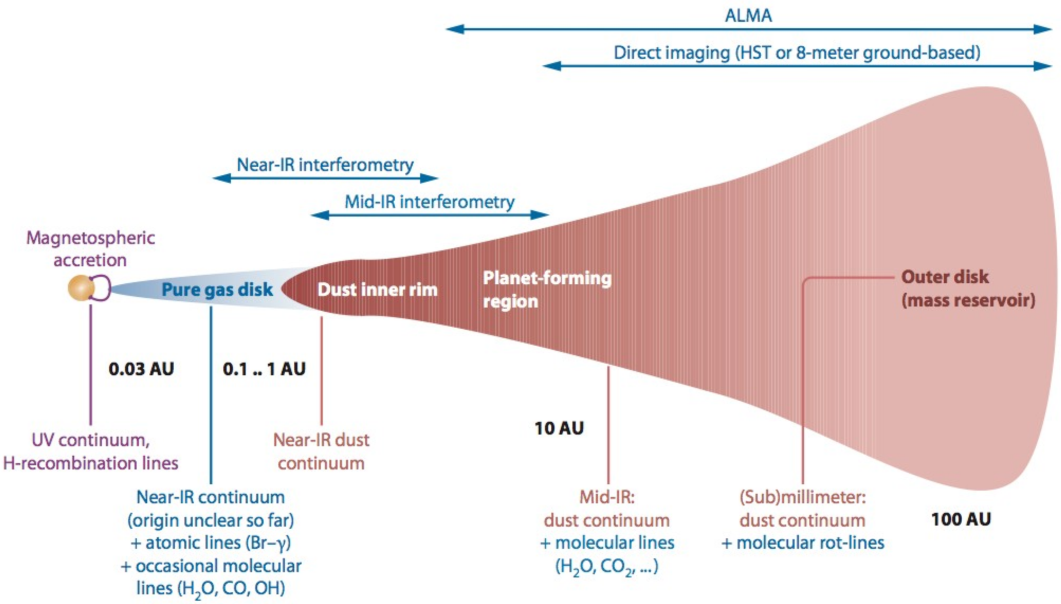
Gravitation „wins“ Magnetic field, Shock wave Protostar

← FREE GAS → ← NO FREE GAS →

Protostars



Isella (2006)

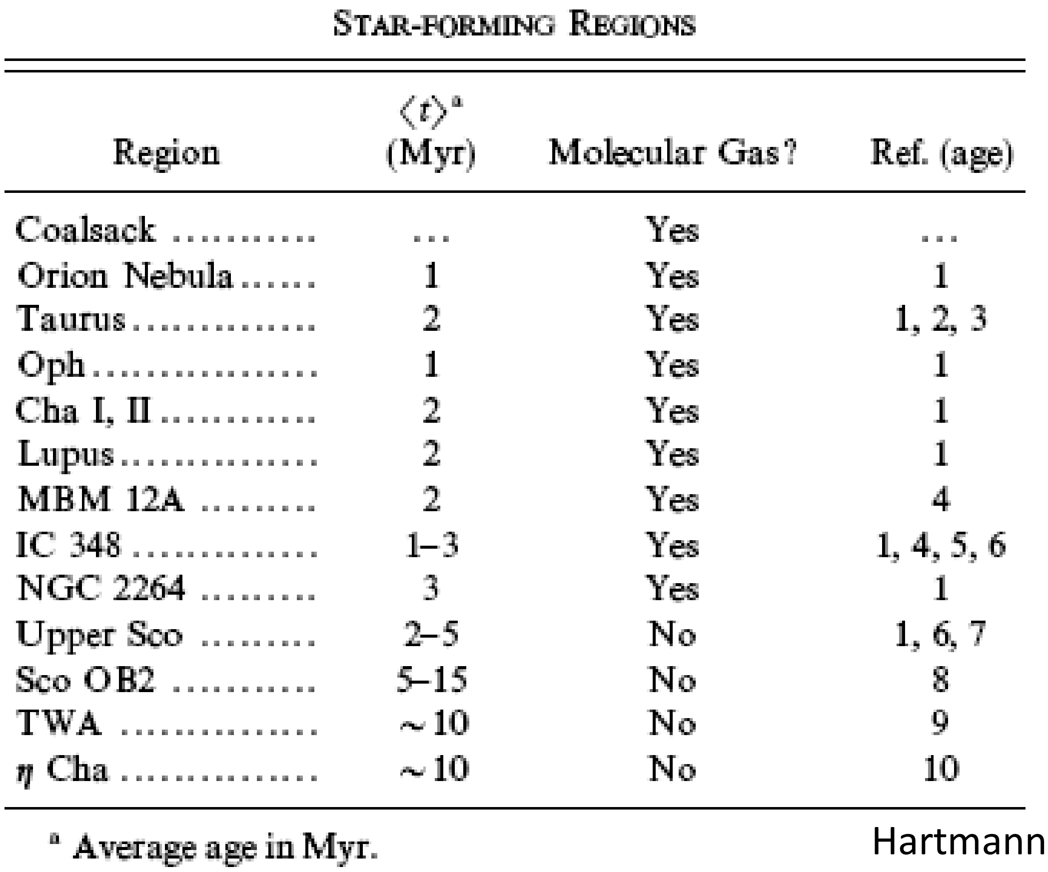


Dullemond & Monnier (2010)

- Relatively young field of study (advent of advanced infrared mission and radio astronomy)
- IR photometry can be used to distinguish between the protostar classes
- High mass/low mass stars form the same way?
- Various timescales for different stellar masses

Star formation

- The detection of free Gas in a Star Cluster is an excellent indicator for the time scale of continuous stellar formation



STAR-FORMING REGIONS

Region	$\langle t \rangle$[a] (Myr)	Molecular Gas?	Ref. (age)
Coalsack	...	Yes	...
Orion Nebula	1	Yes	1
Taurus	2	Yes	1, 2, 3
Oph	1	Yes	1
Cha I, II	2	Yes	1
Lupus	2	Yes	1
MBM 12A	2	Yes	4
IC 348	1–3	Yes	1, 4, 5, 6
NGC 2264	3	Yes	1
Upper Sco	2–5	No	1, 6, 7
Sco OB2	5–15	No	8
TWA	~10	No	9
η Cha	~10	No	10

[a] Average age in Myr.

Star formation lasts 3 to 4 Myrs and is continuous

This is also the "intrinsic" error of an age determination

Hartmann et al., 2001, ApJ, 562, 852

Numerical simulation of star formation in Giant Molecular Clouds

- Hypothesis: the formation of all members of a star cluster is continuous for 3 to 4 Myrs within one GMCs
- Is this a realistic approach?
- Is it possible to simulation the formation of star clusters and compare the results with observational data within the solar vicinity?

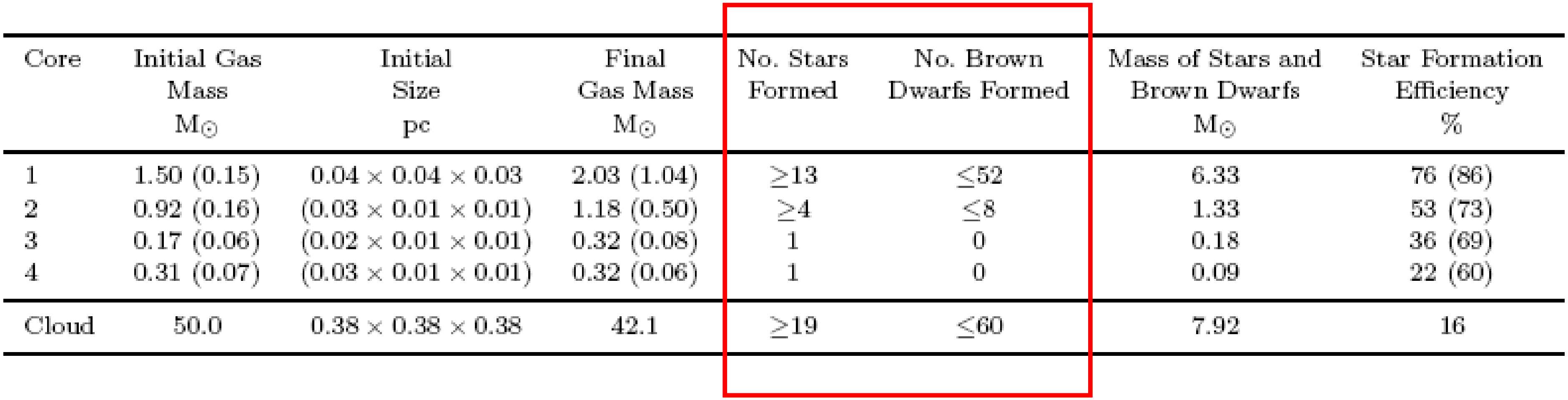
Numerical simulation of star formation in Giant Molecular Clouds

- Detailed paper by Bate & Bonnell, 2005, MNRAS, 356, 1201
- Basis: Orion Nebula and Taurus star forming region
- “Complete” astrophysical numerical simulation including Shock Waves, dynamical parameters and 3D-Hydrodynamics, Jeans Mass < 1 M(sun)
- The numerical simulations are astonishing close to the observations

Numerical simulation of star formation in Giant Molecular Clouds

Input parameter:

1. Mass (GMC) = 50 M(sun), limited by CPU time
2. Diameter = 0.375 pc, limited by CPU time
3. Time for the gravitational collapse: 19000 years
4. Random turbulence field with a 3D Gaussian distribution

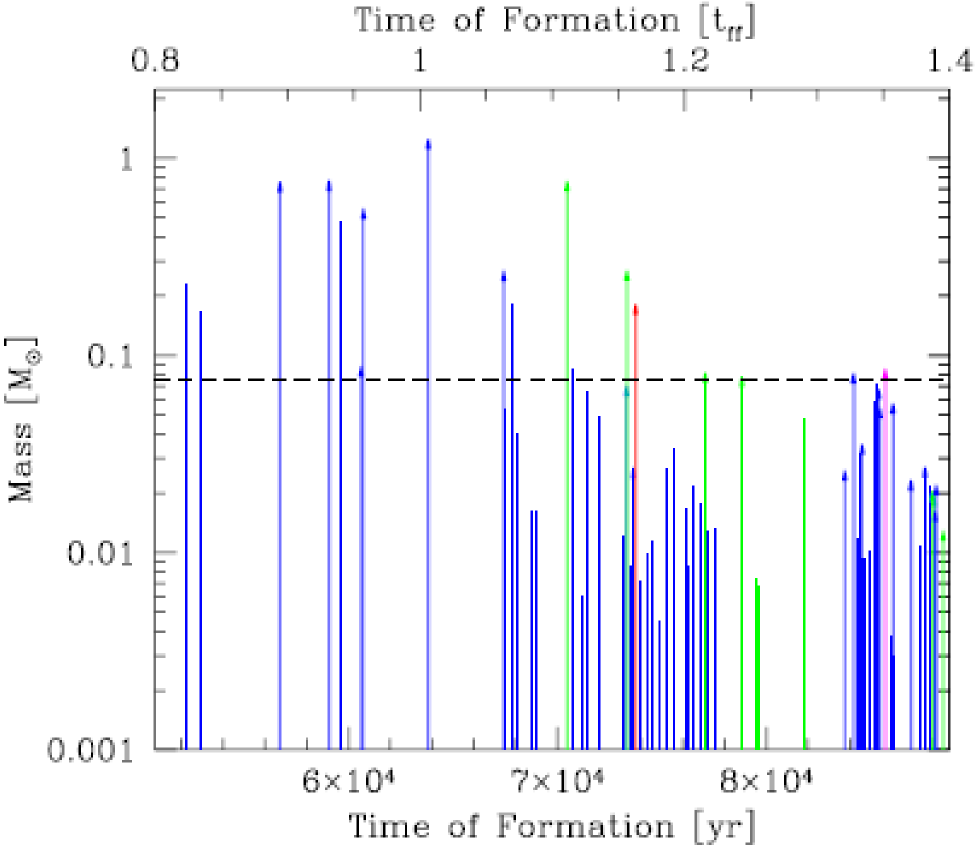


Core	Initial Gas Mass $M_\odot$	Initial Size pc	Final Gas Mass $M_\odot$	No. Stars Formed	No. Brown Dwarfs Formed	Mass of Stars and Brown Dwarfs $M_\odot$	Star Formation Efficiency %
1	1.50 (0.15)	$0.04 \times 0.04 \times 0.03$	2.03 (1.04)	$\geq$13	$\leq$52	6.33	76 (86)
2	0.92 (0.16)	$(0.03 \times 0.01 \times 0.01)$	1.18 (0.50)	$\geq$4	$\leq$8	1.33	53 (73)
3	0.17 (0.06)	$(0.02 \times 0.01 \times 0.01)$	0.32 (0.08)	1	0	0.18	36 (69)
4	0.31 (0.07)	$(0.03 \times 0.01 \times 0.01)$	0.32 (0.06)	1	0	0.09	22 (60)
Cloud	50.0	$0.38 \times 0.38 \times 0.38$	42.1	$\geq$19	$\leq$60	7.92	16

„Stars“: Mass > 0.084 M(sun)
Brown Dwarfs: Mass < 0.084 M(sun), no Hydrogen burning

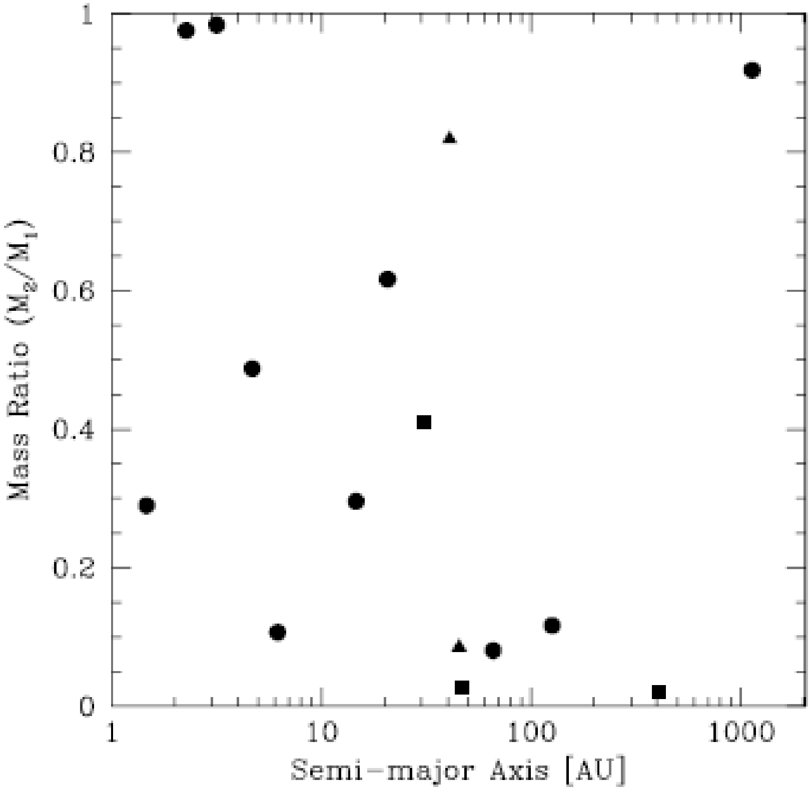
More low mass stars formed due to the IMF

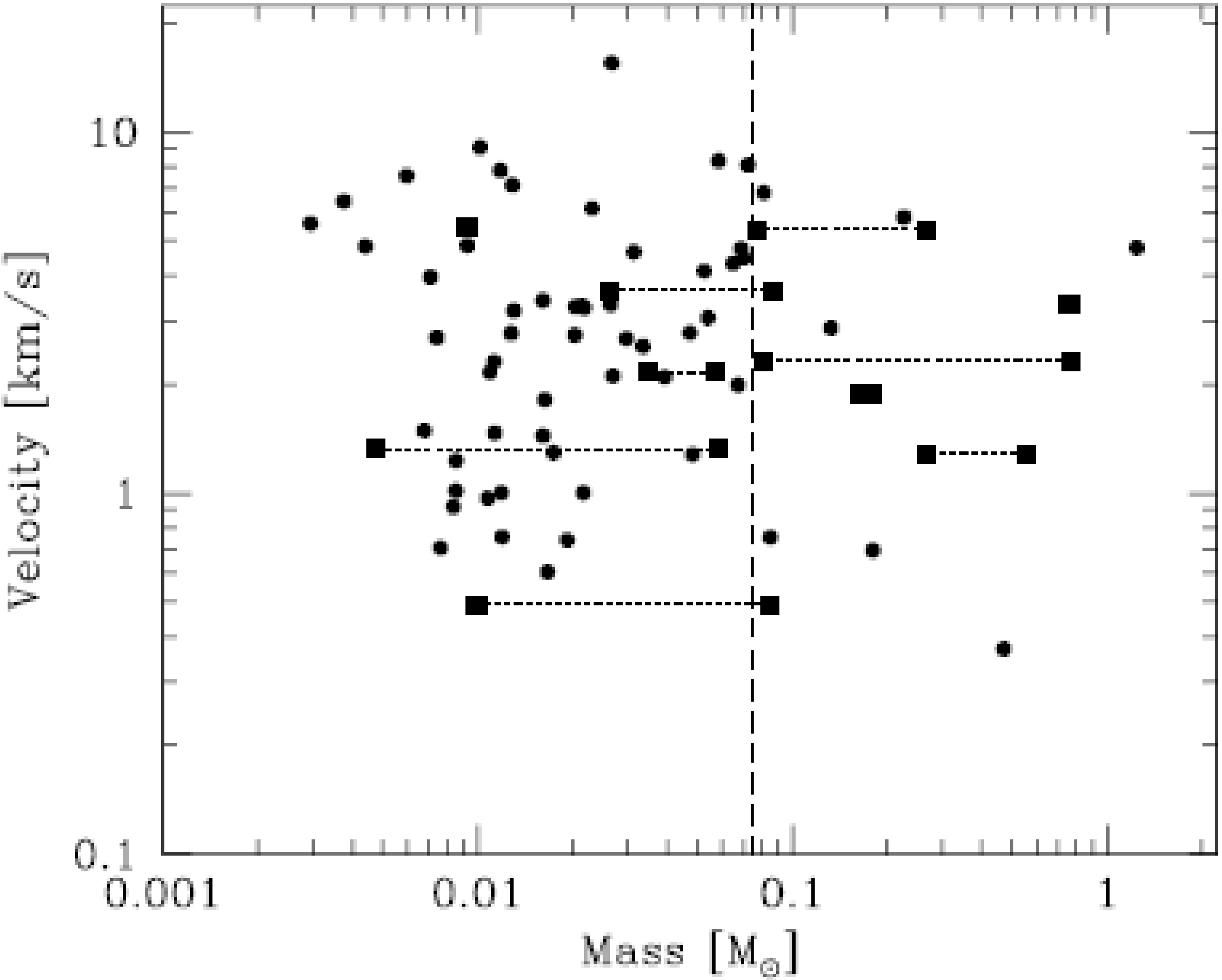
For star clusters it is essential to know the internal velocity distribution because of their evolution (see later)



Continuous star formation in time

The formation of Binary systems





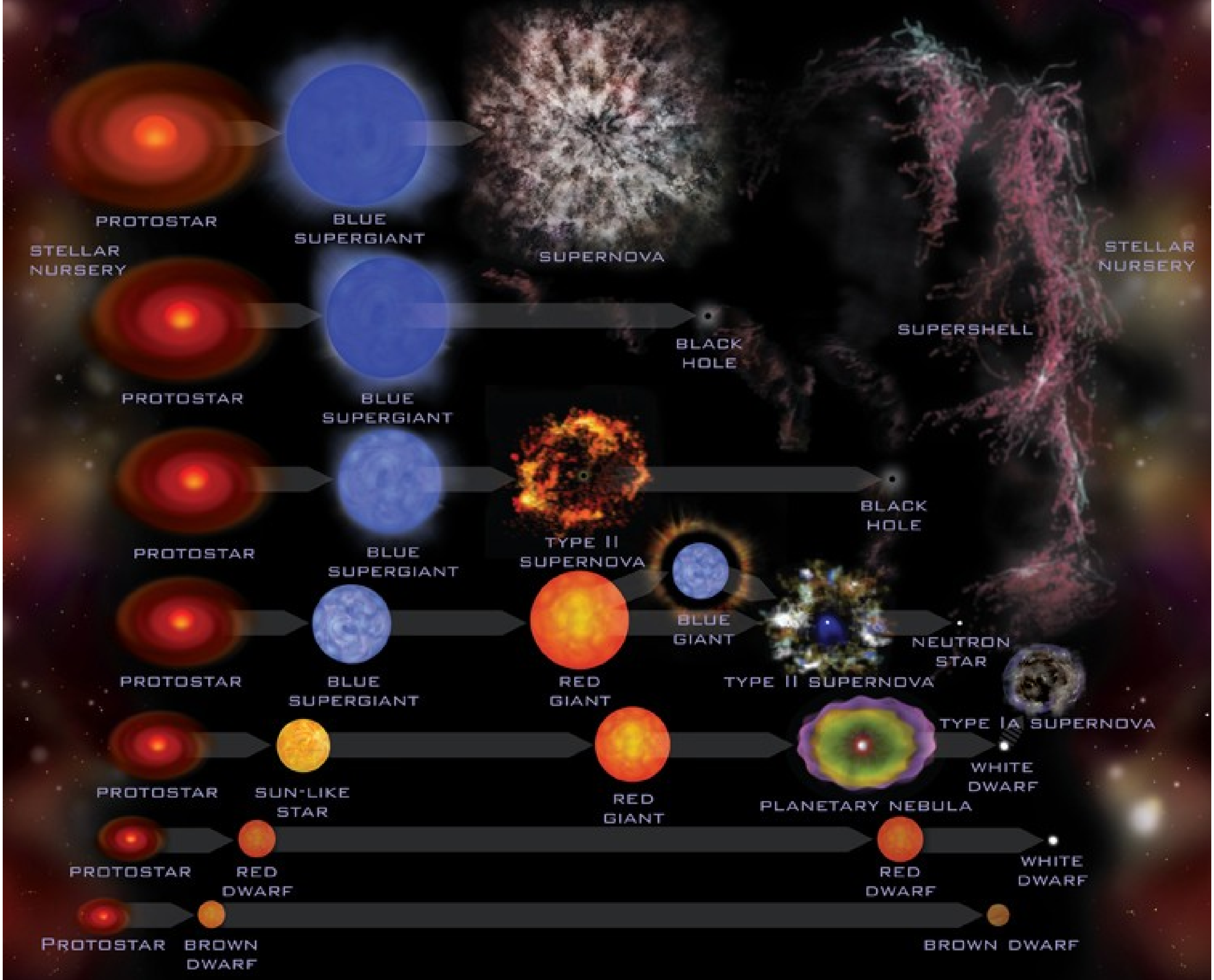
Binaries are connected with a line

The rms velocity dispersion of the simulations is 4.3 km s^{-1}
Such observational data for d > 500 pc are still not available => Gaia satellite mission

Evolution of Star Clusters

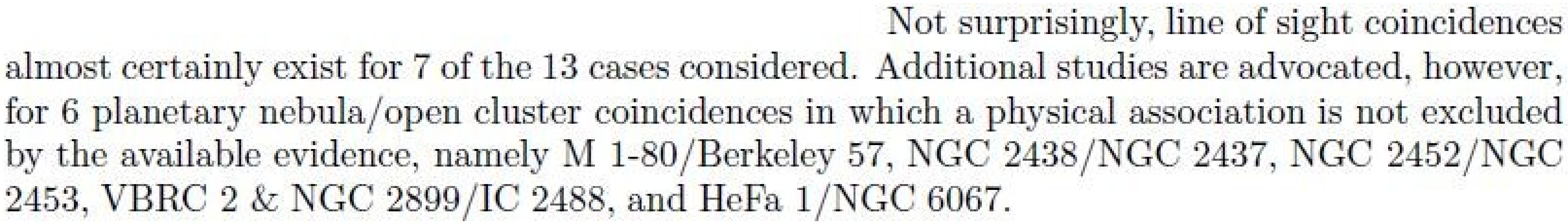
- Star Clusters form with the following characteristics
  1. **Total Mass**: IMF
  2. **Metallicity**
  3. **Kinematics of the Cluster center**: location within the Galaxy
  4. **Internal velocity dispersion**
- How does a Star Cluster evolve with these starting parameters?

- Each member (= star) evolve “as an individual”, some important topics
    1. Binary Evolution
    2. Mass Loss (hot stars)
    3. AGB Evolution
    4. Planetary Nebula (cool stars)
    5. Supernovae explosions
- In Star Clusters, collisions are very uncommon (see later), almost no new multiple (binary) systems form during the later evolution
- Star Clusters, normally, follow Galactic Rotation

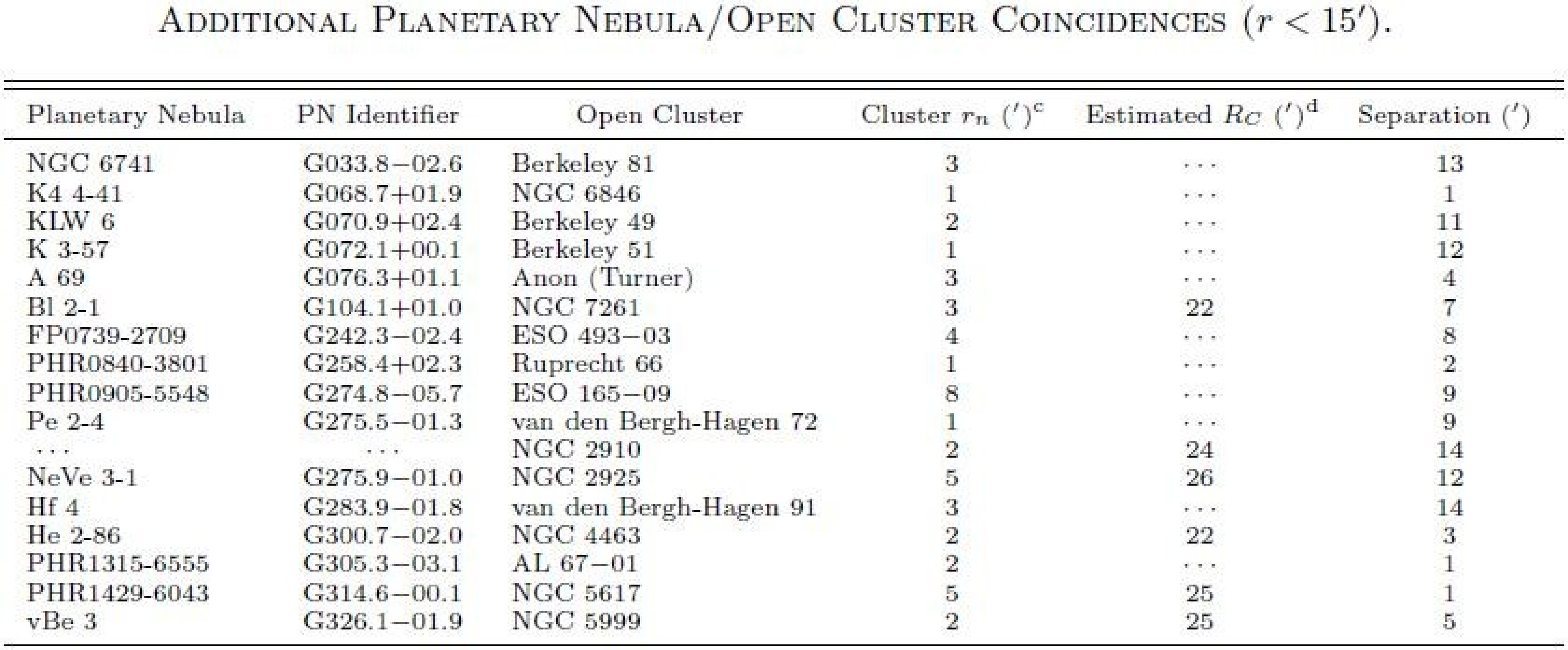

PROTOSTAR
BLUE SUPERGIANT
SUPERNOVA
STELLAR NURSERY
STELLAR NURSERY
PROTOSTAR
BLUE SUPERGIANT
BLACK HOLE
SUPERSHELL
PROTOSTAR
BLUE SUPERGIANT
TYPE II SUPERNOVA
BLACK HOLE
PROTOSTAR
BLUE SUPERGIANT
RED GIANT
BLUE GIANT
TYPE II SUPERNOVA
NEUTRON STAR
TYPE IA SUPERNOVA
PROTOSTAR
SUN-LIKE STAR
RED GIANT
PLANETARY NEBULA
WHITE DWARF
PROTOSTAR
RED DWARF
RED DWARF
WHITE DWARF
PROTOSTAR
BROWN DWARF
BROWN DWARF

Planetary Nebulae

Majaess et al., 2007, PASP, 119, 1349



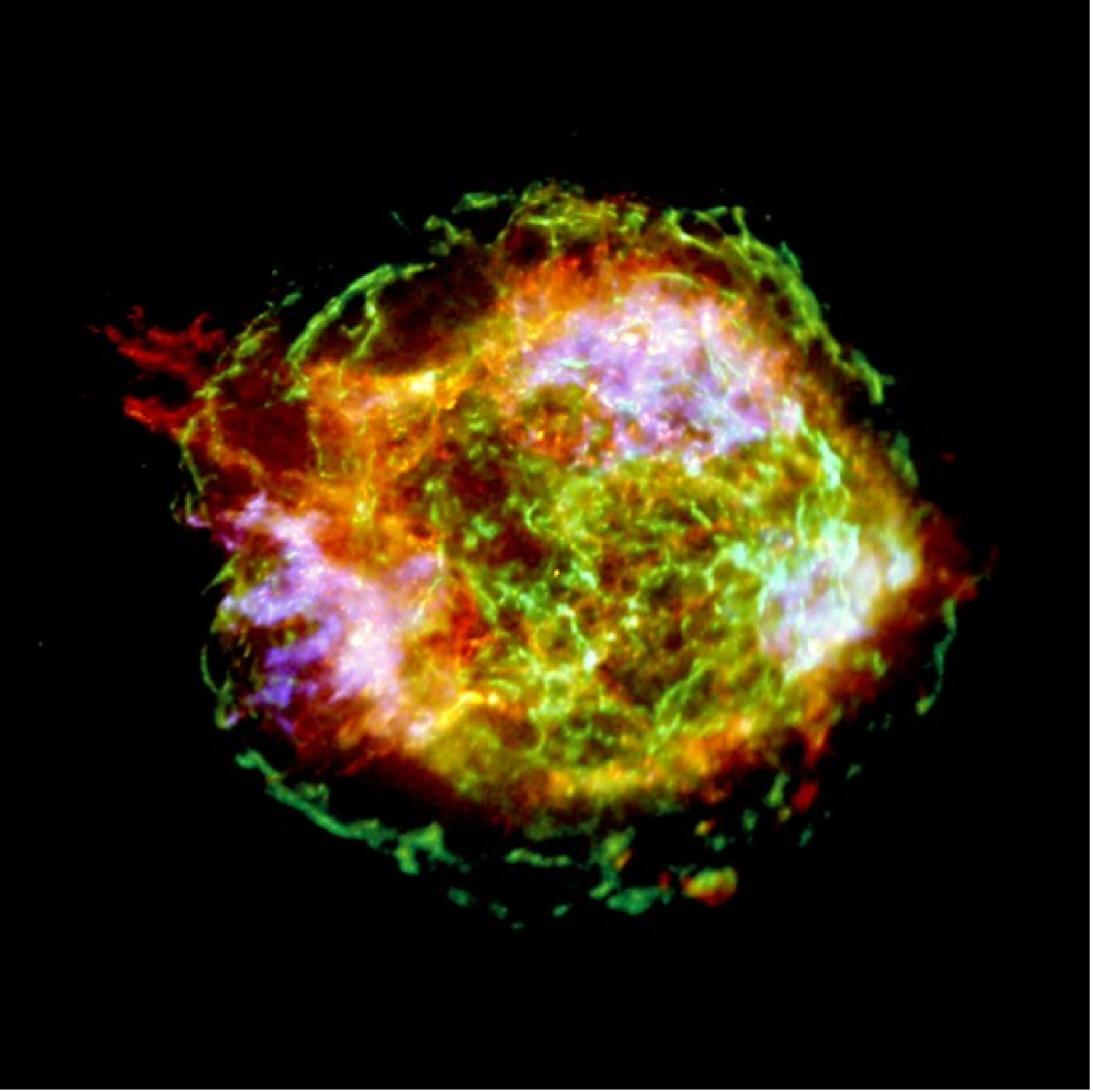
Not surprisingly, line of sight coincidences almost certainly exist for 7 of the 13 cases considered. Additional studies are advocated, however, for 6 planetary nebula/open cluster coincidences in which a physical association is not excluded by the available evidence, namely M 1-80/Berkeley 57, NGC 2438/NGC 2437, NGC 2452/NGC 2453, VBRC 2 & NGC 2899/IC 2488, and HeFa 1/NGC 6067.



ADDITIONAL PLANETARY NEBULA/OPEN CLUSTER COINCIDENCES ($r < 15'$).

Planetary Nebula	PN Identifier	Open Cluster	Cluster r_n (′)[c]	Estimated R_C (′)[d]	Separation (′)
NGC 6741	G033.8−02.6	Berkeley 81	3	...	13
K4 4-41	G068.7+01.9	NGC 6846	1	...	1
KLW 6	G070.9+02.4	Berkeley 49	2	...	11
K 3-57	G072.1+00.1	Berkeley 51	1	...	12
A 69	G076.3+01.1	Anon (Turner)	3	...	4
Bl 2-1	G104.1+01.0	NGC 7261	3	22	7
FP0739-2709	G242.3−02.4	ESO 493−03	4	...	8
PHR0840-3801	G258.4+02.3	Ruprecht 66	1	...	2
PHR0905-5548	G274.8−05.7	ESO 165−09	8	...	9
Pe 2-4	G275.5−01.3	van den Bergh-Hagen 72	1	...	9
...	...	NGC 2910	2	24	14
NeVe 3-1	G275.9−01.0	NGC 2925	5	26	12
Hf 4	G283.9−01.8	van den Bergh-Hagen 91	3	...	14
He 2-86	G300.7−02.0	NGC 4463	2	22	3
PHR1315-6555	G305.3−03.1	AL 67−01	2	...	1
PHR1429-6043	G314.6−00.1	NGC 5617	5	25	1
vBe 3	G326.1−01.9	NGC 5999	2	25	5

PNs exist in Open Clusters



Important topic of how SN explosions affect the cluster evolution

Shockwaves
Mass flow

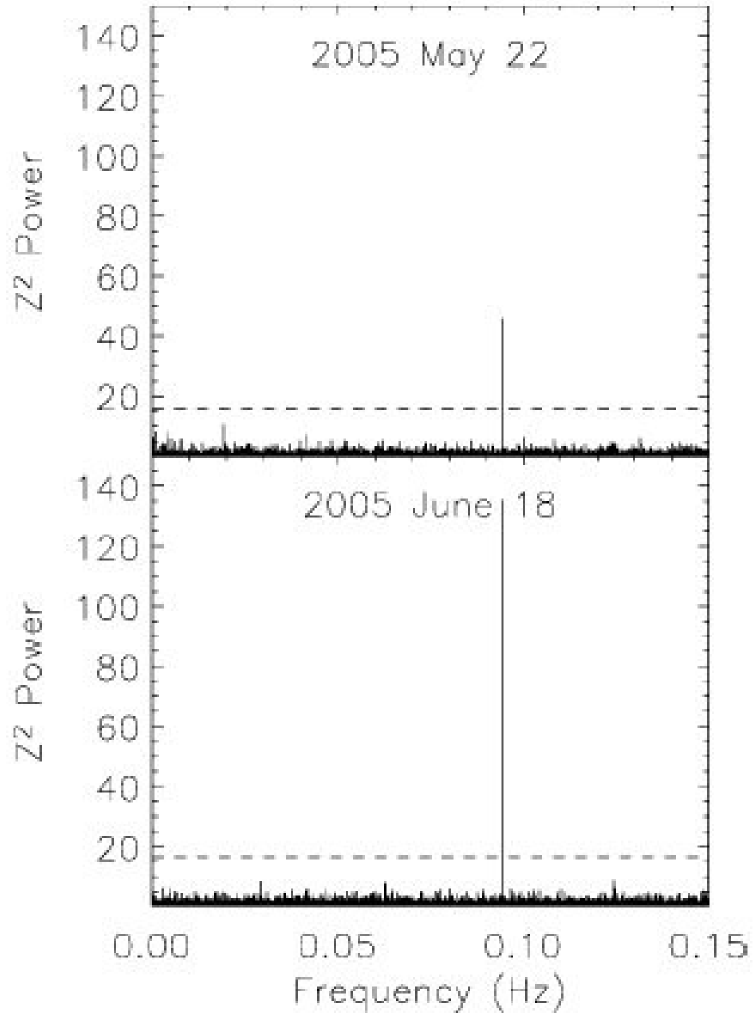
Statistically, SN explosions are rather common

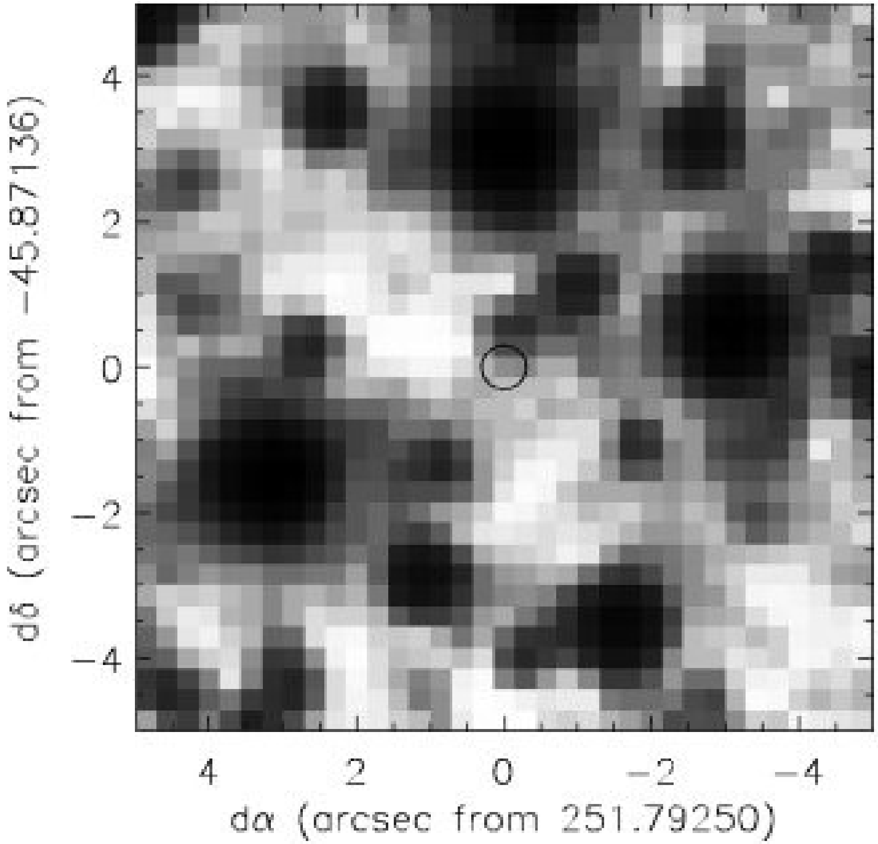
From Chandra, Diameter 4 pc

SN Remnants

- Catalogue of galactic SNRs: http://www.mrao.cam.ac.uk/surveys/snrs/
- 274 entries
- Complete list of papers for Open Clusters
  1. Pauls, 1977, A&A, 59, L13: NGC 559?
  2. Kumar, 1978, ApJ, 219, L13: **Tr 18** and 21?
  3. Peterson et al., 1988, MNRAS, 235, 1439: Lynga 1, Pismis 20, Stock 14, and Trumpler 21, none conclusive

Muno et al., 2006, ApJ, 636, L41: **Westerlund 1**
d = 5200 pc, log t < 6.4

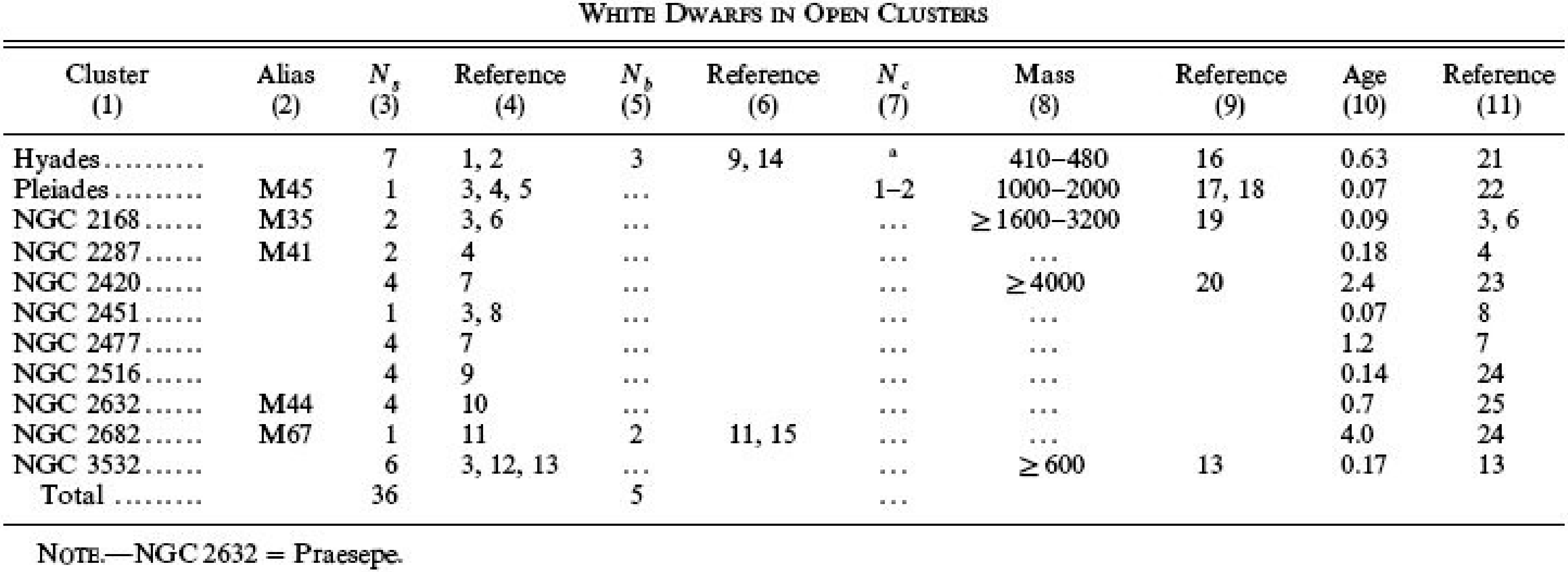




Pulsar, V fainter than 25th mag

- White Dwarfs were detected in Open Clusters
- The number is compatible with a common stellar evolution scenario, but the membership determination is very difficult
- The absolute magnitude of WDs is about 10 magnitudes fainter than the corresponding Main Sequence

von Hippel, 1998, AJ, 115, 1536



White Dwarfs in Open Clusters

Cluster (1)	Alias (2)	N_s (3)	Reference (4)	N_b (5)	Reference (6)	N_c (7)	Mass (8)	Reference (9)	Age (10)	Reference (11)
Hyades..........		7	1, 2	3	9, 14	[a]	410–480	16	0.63	21
Pleiades.........	M45	1	3, 4, 5	...		1–2	1000–2000	17, 18	0.07	22
NGC 2168......	M35	2	3, 6	...		...	≥1600–3200	19	0.09	3, 6
NGC 2287......	M41	2	4	...		...	...		0.18	4
NGC 2420......		4	7	...		...	≥4000	20	2.4	23
NGC 2451......		1	3, 8	...		...	...		0.07	8
NGC 2477......		4	7	...		...	...		1.2	7
NGC 2516......		4	9	...		...	...		0.14	24
NGC 2632......	M44	4	10	...		...	...		0.7	25
NGC 2682......	M67	1	11	2	11, 15	...	...		4.0	24
NGC 3532......		6	3, 12, 13	...		...	≥600	13	0.17	13
Total		36		5		...				

Note.—NGC 2632 = Praesepe.

Single Multiple

In total, 41 WDs until 1998 found, no firm improvement after that

Why do Star Clusters dissipate?

Virial Theorem: $2E_{kin} = -\Omega$

Kinetic Energy: $2E_{kin} = n \cdot m_i \cdot \bar{v}^2 = M \cdot \bar{v}^2$

$\bar{v}\ldots$ mean v of the members relative to the cluster center

Potential Energy:

$$\Omega \quad -\frac{1}{2} \cdot \frac{G \cdot M^2}{\bar{R}^2}$$

yielding:

$$\bar{v}^2 = \frac{G \cdot M}{2\bar{R}^2}$$

Escape Velocity: $\bar{v}_\infty^2 = \cdot \bar{v}^2$

Collisions: $t_{coll} \approx \dfrac{1}{\rho\ \ \sigma\ \ \Delta}$

Density ρ and cross section σ:

$$\rho = \frac{N}{\bar{R}^3} \qquad \sigma = 4\pi\ \ R_*^2 \Rightarrow \quad {}_{oll} = \frac{\bar{R}^3}{4\pi\ \ N \cdot R_*^2 \cdot \Delta}$$

Example of a typical Open Cluster:

$$N = \ 000, \Delta v = \ 0\mathrm{kms}^{-1}, \mathrm{R}_* = \ .5 R_{Sun}, \bar{\mathrm{R}} = \ \mathrm{pc}$$

$t_{coll} = 10^{25}\mathrm{s}$ => Collisions play no role

Even in the most inner core parts, collisions are highly improper, but could occur

Conclusions:

1. Binary and Multiple systems are not results of collisions in later stages but form already at the very beginning
2. Members do, in general, not escape due to collisions (swing-by effect), but their peculiar velocity component is part of the cluster formation or due to a SN

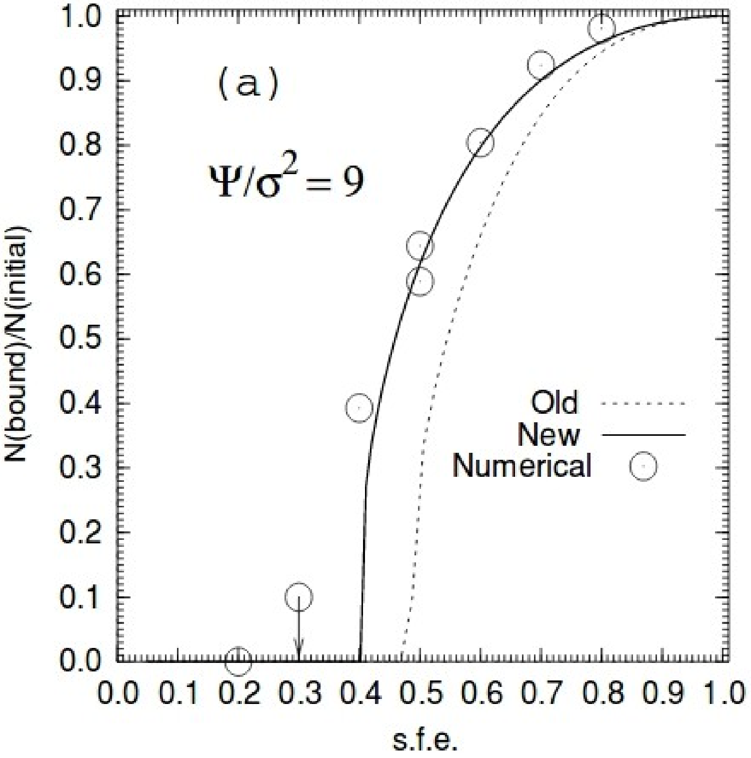
Crossing Time: $t_{cross} = \frac{\overline{R}}{\Delta}$

$$\Delta = 0\mathrm{kms}^{-1} \text{ and } \overline{\mathrm{R}} = 5\mathrm{pc} \Rightarrow {}_{oss} = 4.9 \cdot 10^{8}\, yr$$

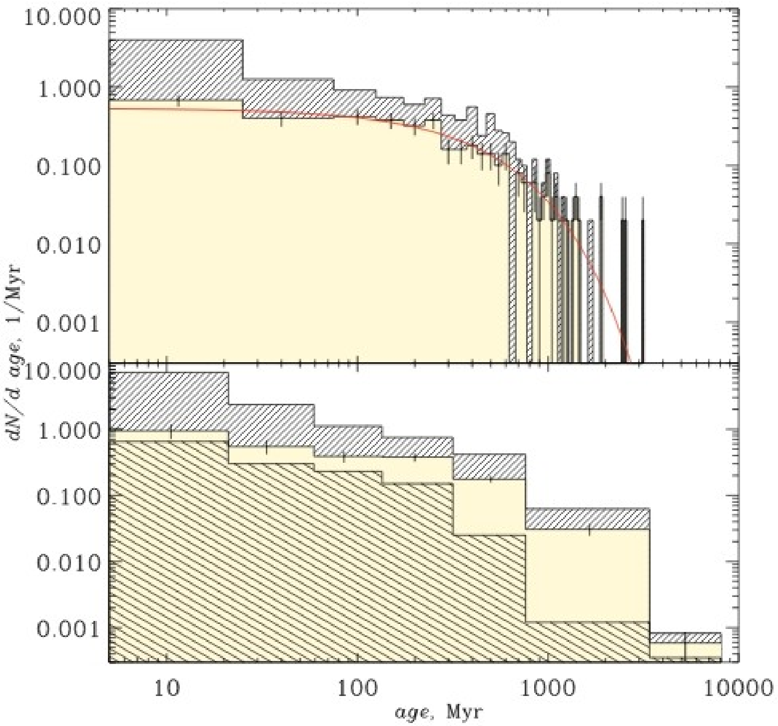
Members can escape from a Star Cluster on a relatively short time scale

Reason: Velocity dispersion caused by the cluster formation and SN events

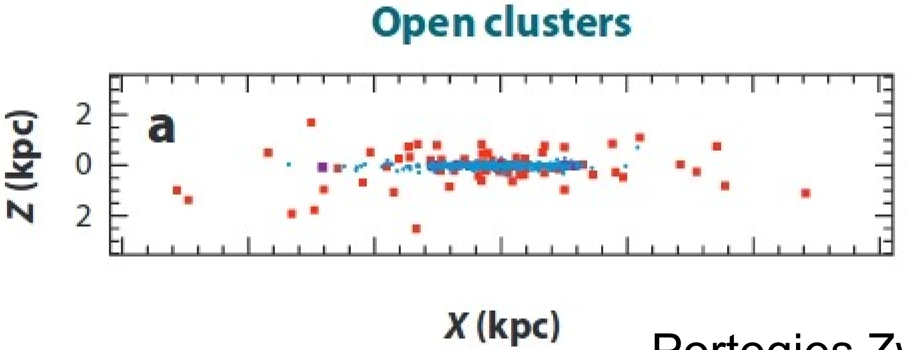
Infant mortality problem



Boily & Kroupa (2003)



Piskunov+ (2006)



Portegies Zwart+ (2010)

Tidal Forces due to Differential Galactic Rotation

Total Mass of the Milky Way: $M_G = 2 \times 10^{11}$ M(Sun)

Gravitational acceleration of the complete star cluster g_G and the individual member g_* :

$$g_G = \frac{G \cdot M_G}{R_{GC}^2} \qquad g_* = \frac{G \cdot M_G}{\left(R_{GC} - \right)^2}$$

The difference of these two values, is the force, of which “the Galaxy” tries to pull away a star from the cluster

$$g_{G,*} = \frac{2\cdot G\cdot M_G\cdot r}{R_{GC}^3} \text{ for } r \ll R_{GC}$$

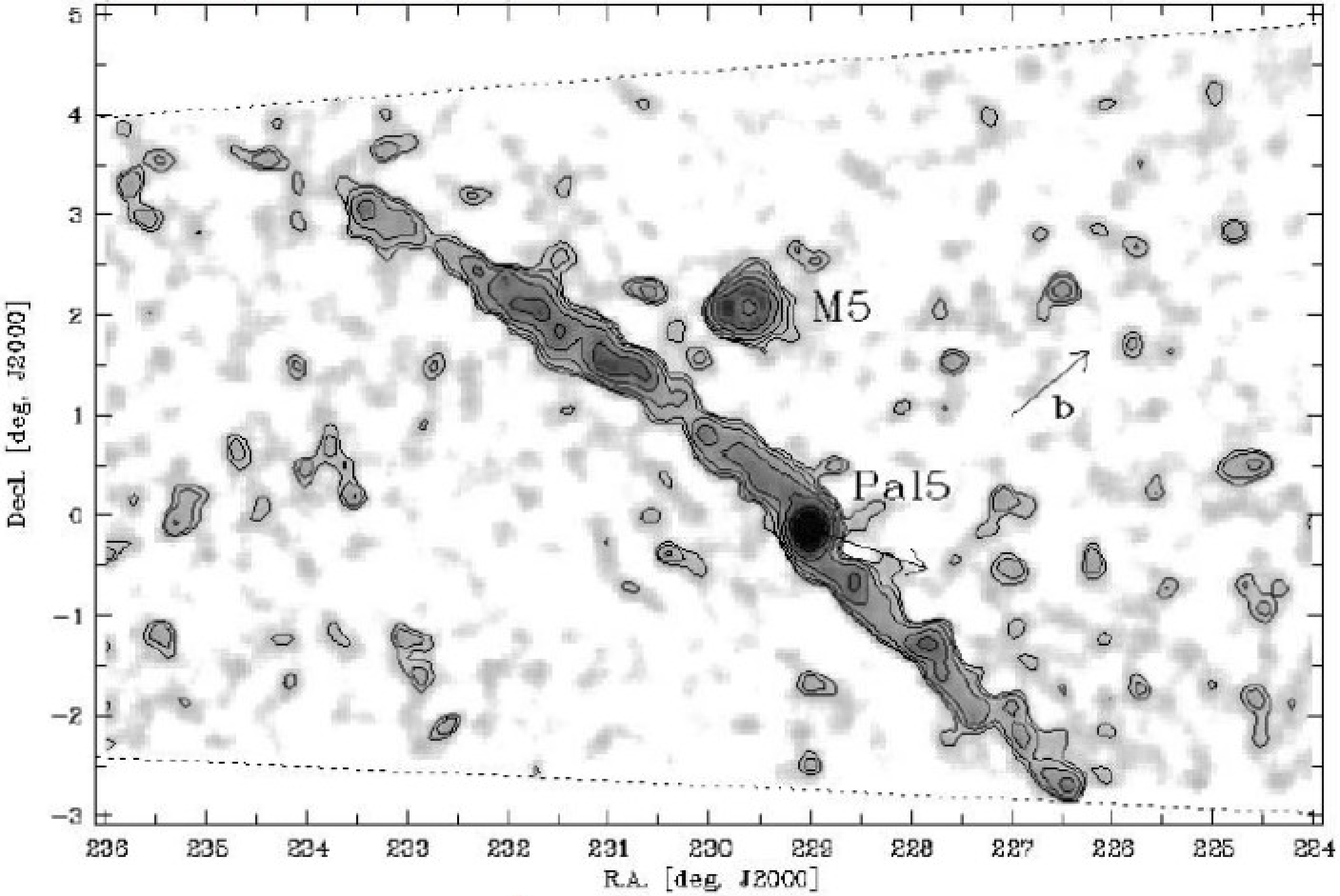
On the other side we have the gravitational force of the open cluster. The stability radius r_S is defined as:

$$\frac{2\cdot G\cdot M_G\cdot r_S}{R_{GC}^3} = \frac{G\cdot M_{OC}}{r_s^2} \Rightarrow r_S = R_{GC}\cdot\left(\frac{M_{OC}}{2M_G}\right)^{1/3}$$

For 1000 M(Sun) => Diameter 20 pc

$$r_S = 0.9\cdot\left(\frac{M_{OC}}{1000}\right)^{1/3} \quad \text{for } R_{GC} = 8\,\text{kpc in } [\text{M}_{\text{Sun}}, \text{pc}]$$

(Odenkirchen et al. 2003)



Summary

- Star Cluster dissipate because of
  1. Differential Galactic Rotation
  2. Internal Velocity Dispersion
  3. Collisions in the first few Myrs
  4. SN Explosions and corresponding Shock Waves
  5. (Collisions with “Field Stars”)
- Explains the existence of Globular Clusters
- Valid for all Spiral Galaxies